

SIMULATING *ASTRO-H* OBSERVATIONS
OF GALAXY CLUSTER GAS MOTIONS:
WHAT WE CAN EXPECT AND
IMPLICATIONS FOR FUTURE MISSIONS

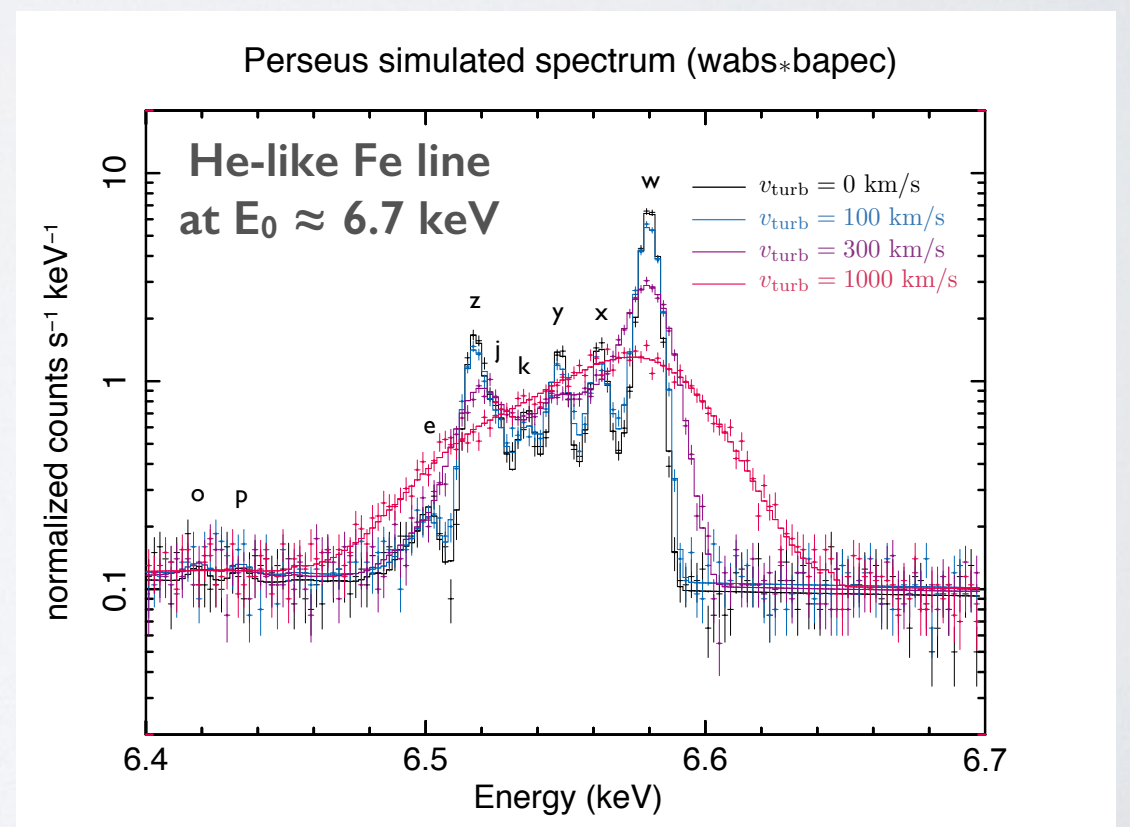
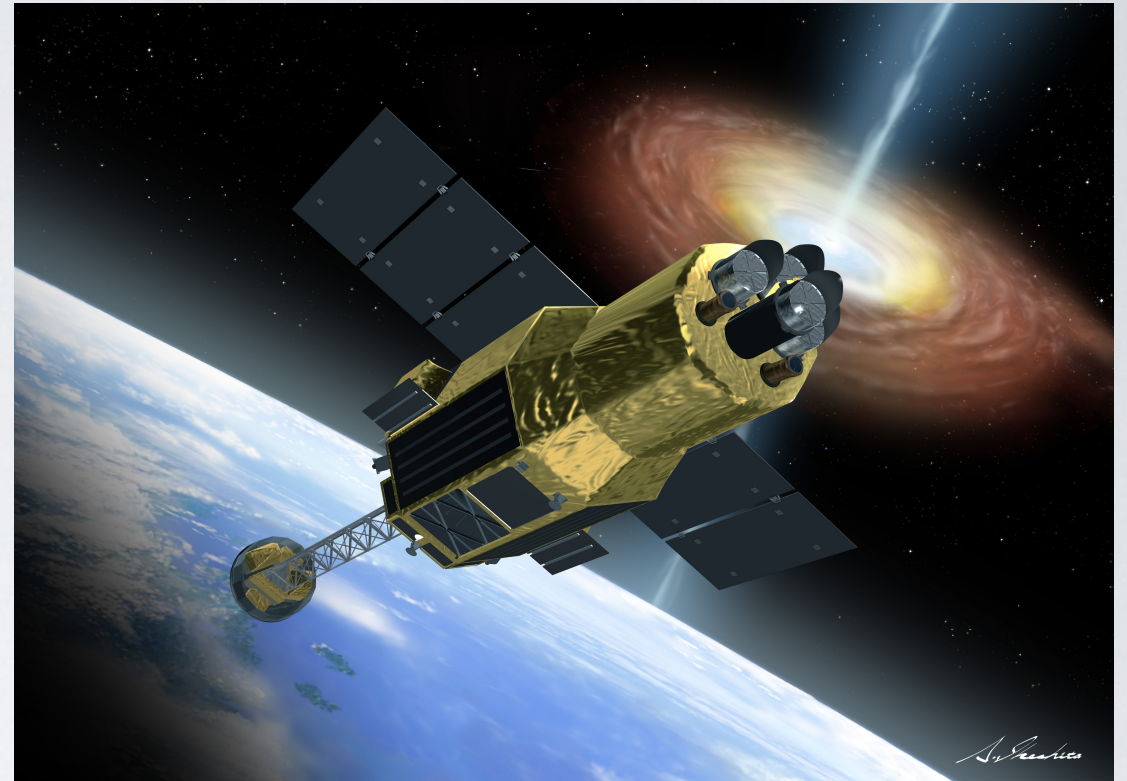
John ZuHone, MIT Kavli Institute

with

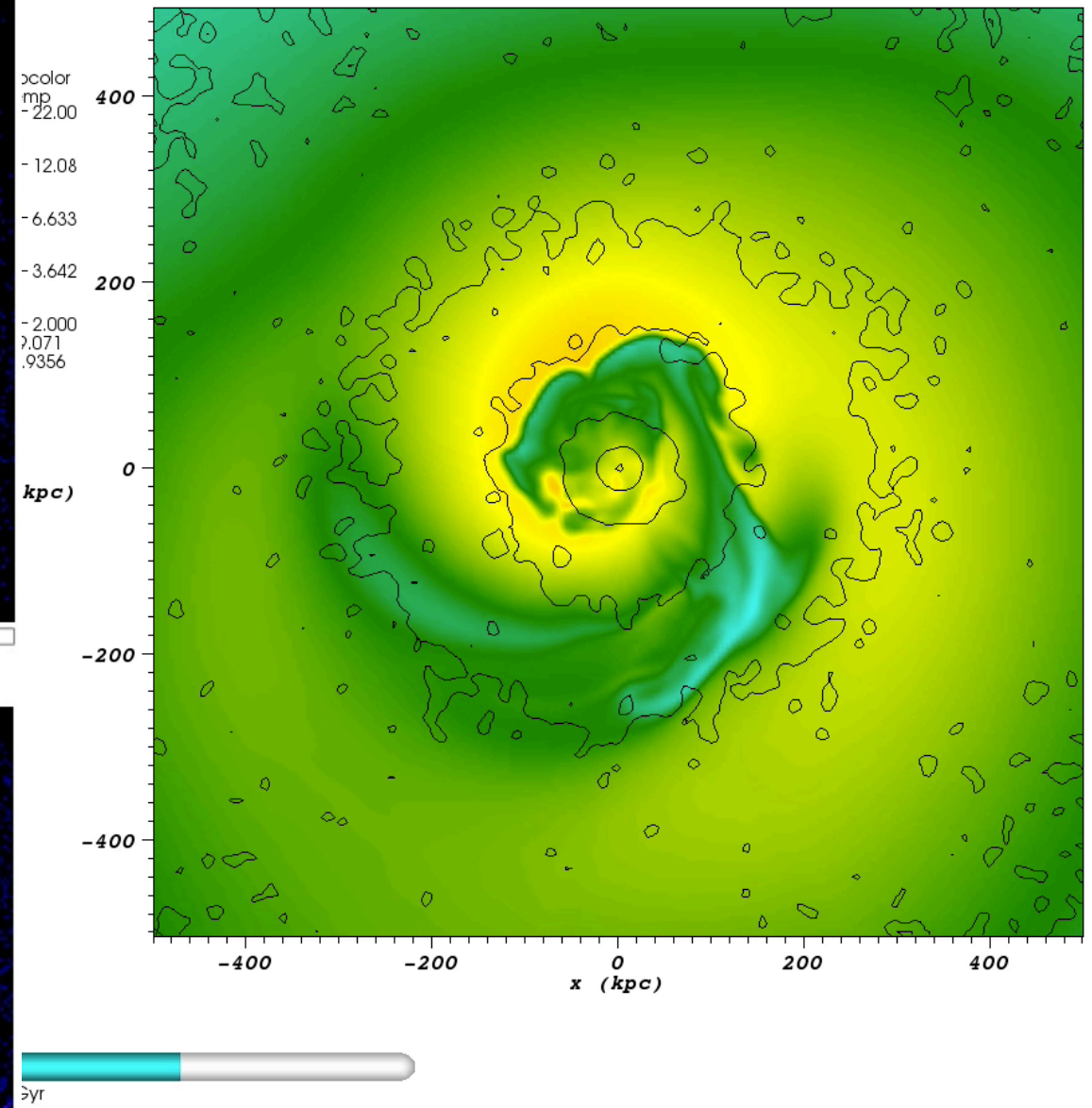
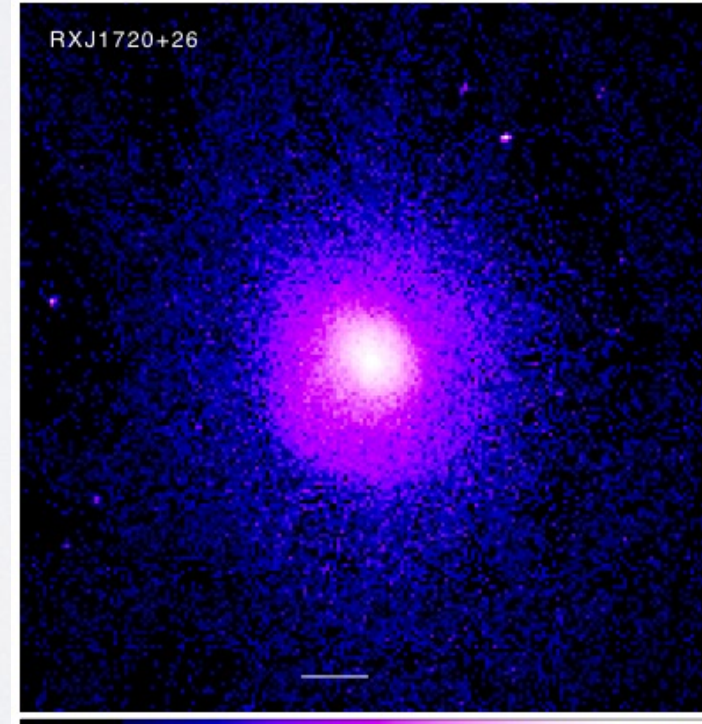
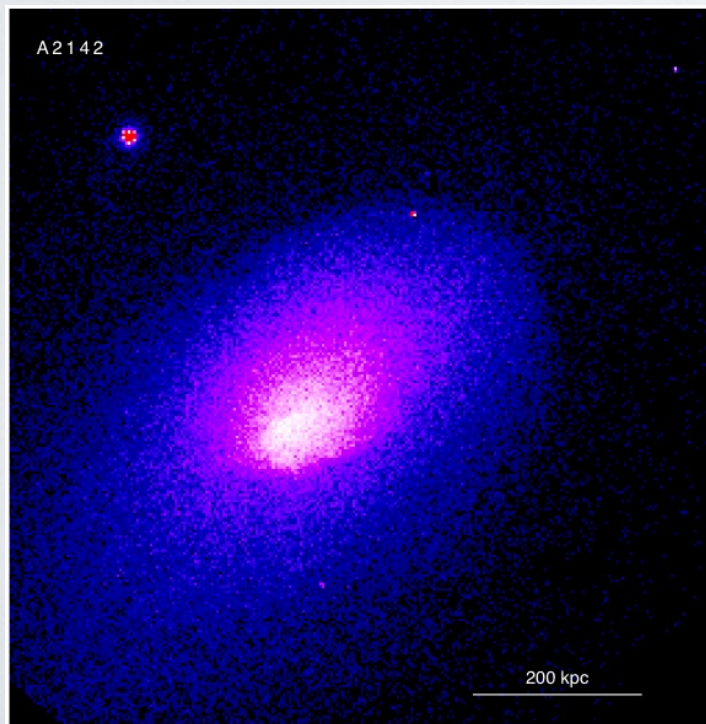
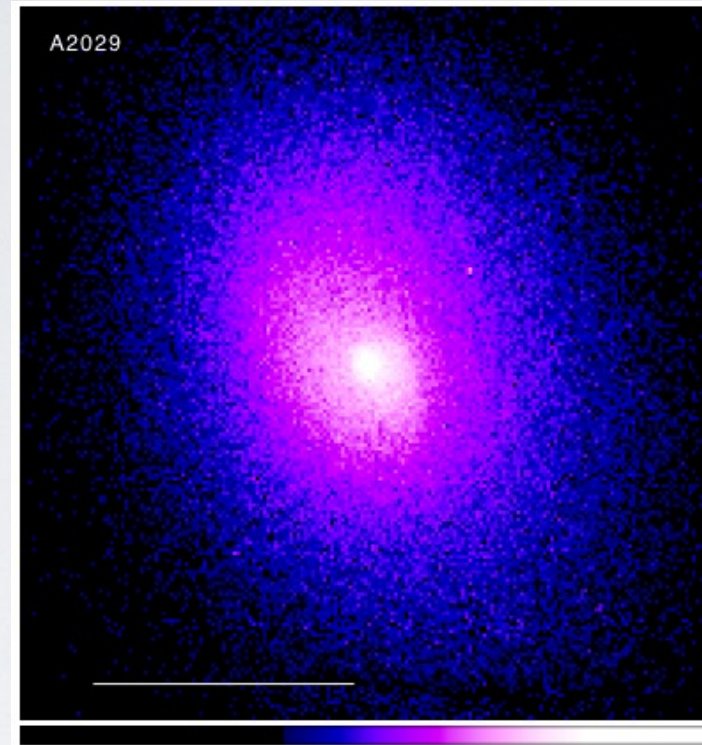
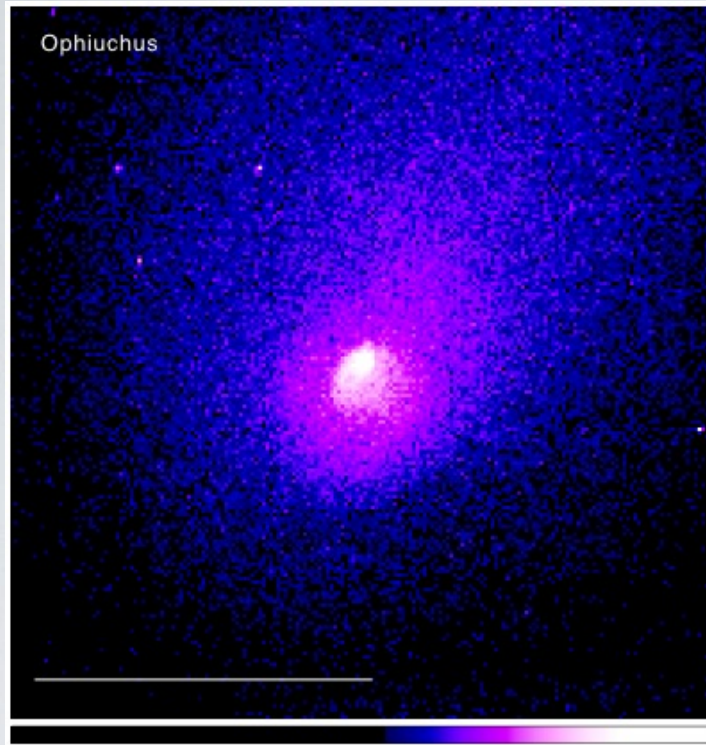
Eric Miller (MIT), Aurora Simionescu (ISAS/JAXA), Mark Bautz (MIT)

ASTRO-H MISSION

- To be launched early next year
- Soft X-ray Spectrometer (SXS):
 - 3'x3' FOV
 - ~5 eV spectral resolution
 - ~1' spatial PSF



CORE GAS SLOSHING



ZuHone, Markevitch,
& Johnson 2010

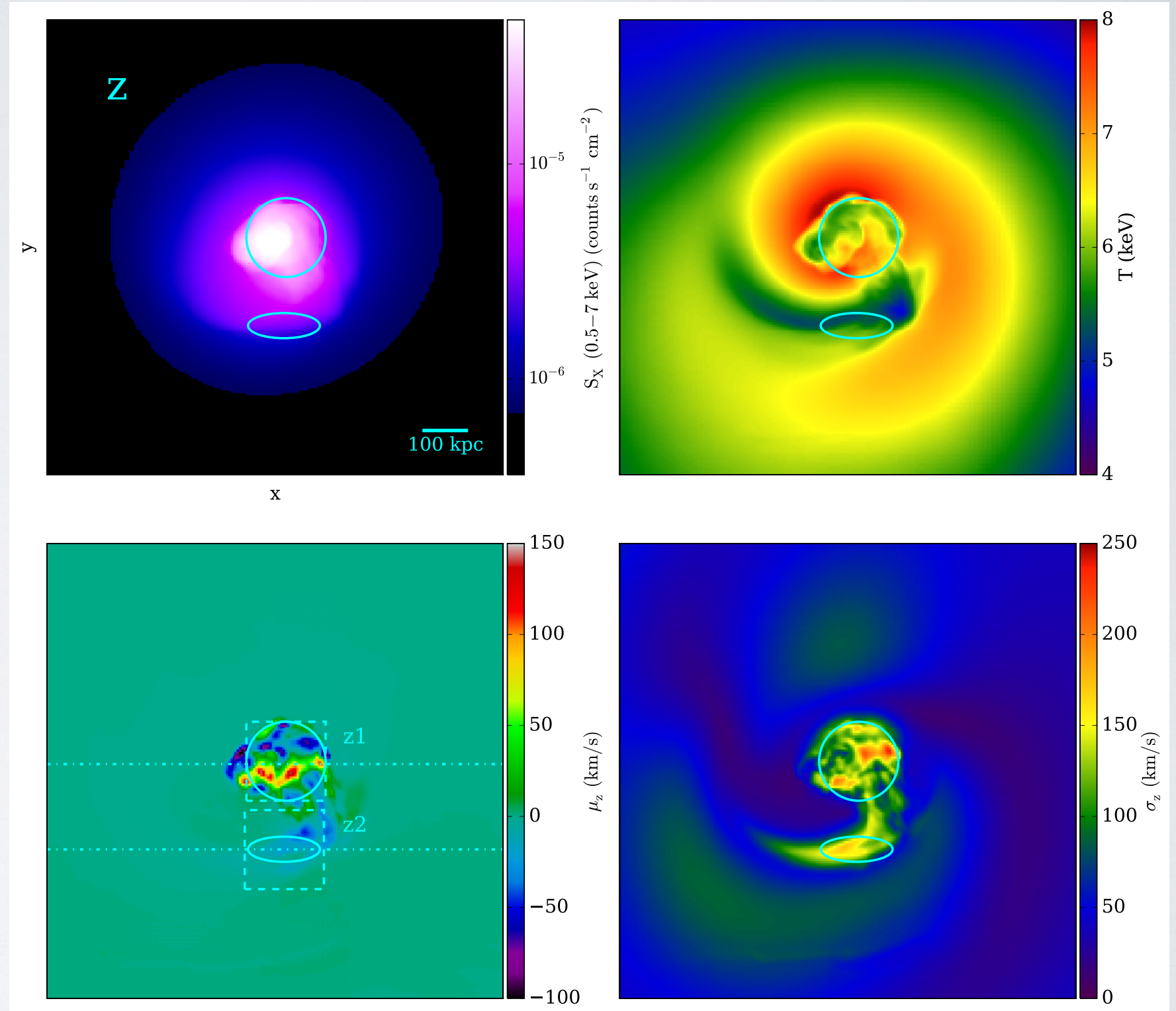
Markevitch & Vikhlinin 2007

- Can we detect these sloshing motions with *Astro-H*?
- If so, what effect will these motions have on the shift and shape of spectral lines?
- Can we use this spectral analysis to tell us something about microphysics?

LOOKING PERP TO PLANE

SB

T



μ

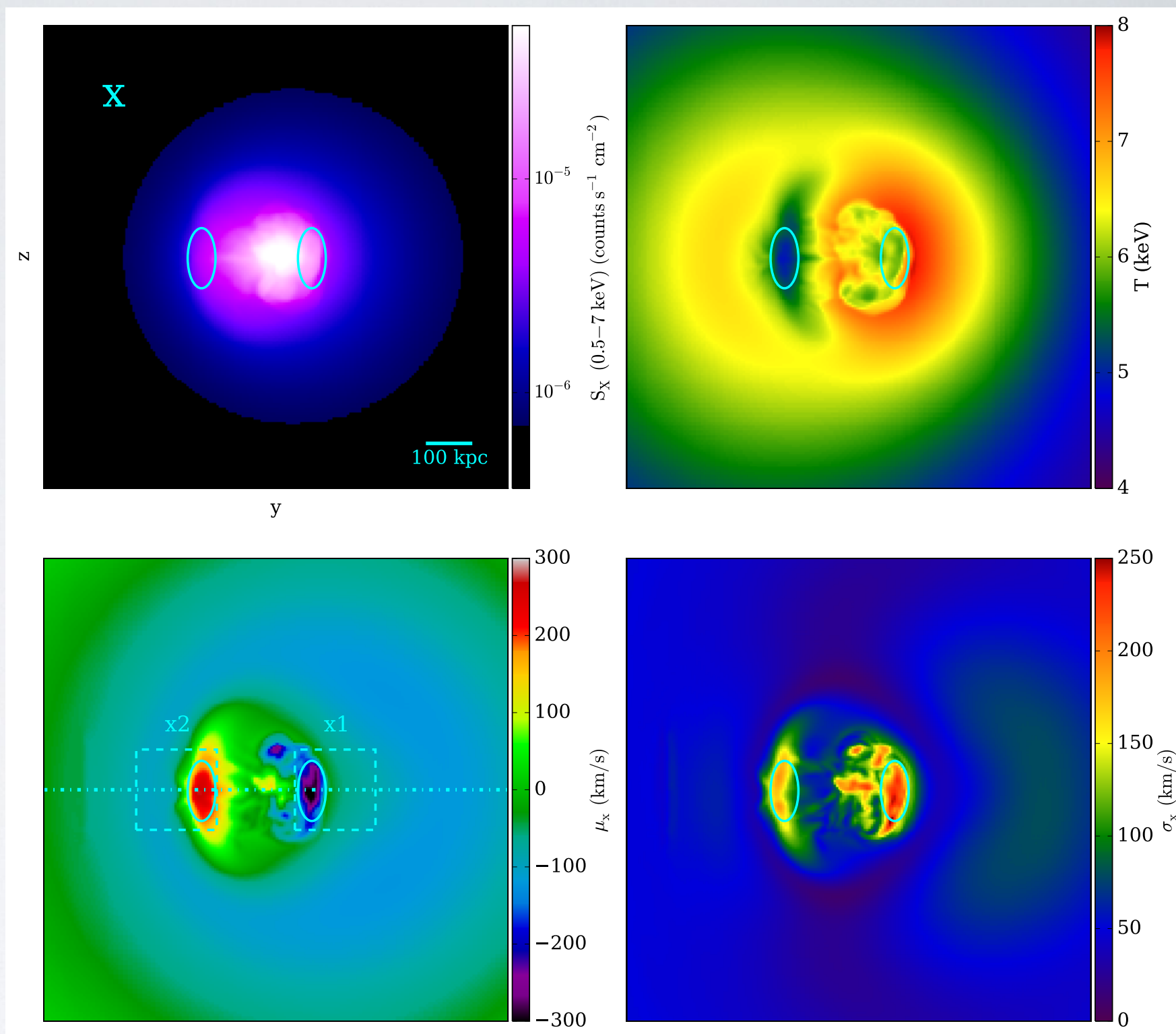
σ

Inviscid

SB

T

LOOKING INTO THE PLANE



μ

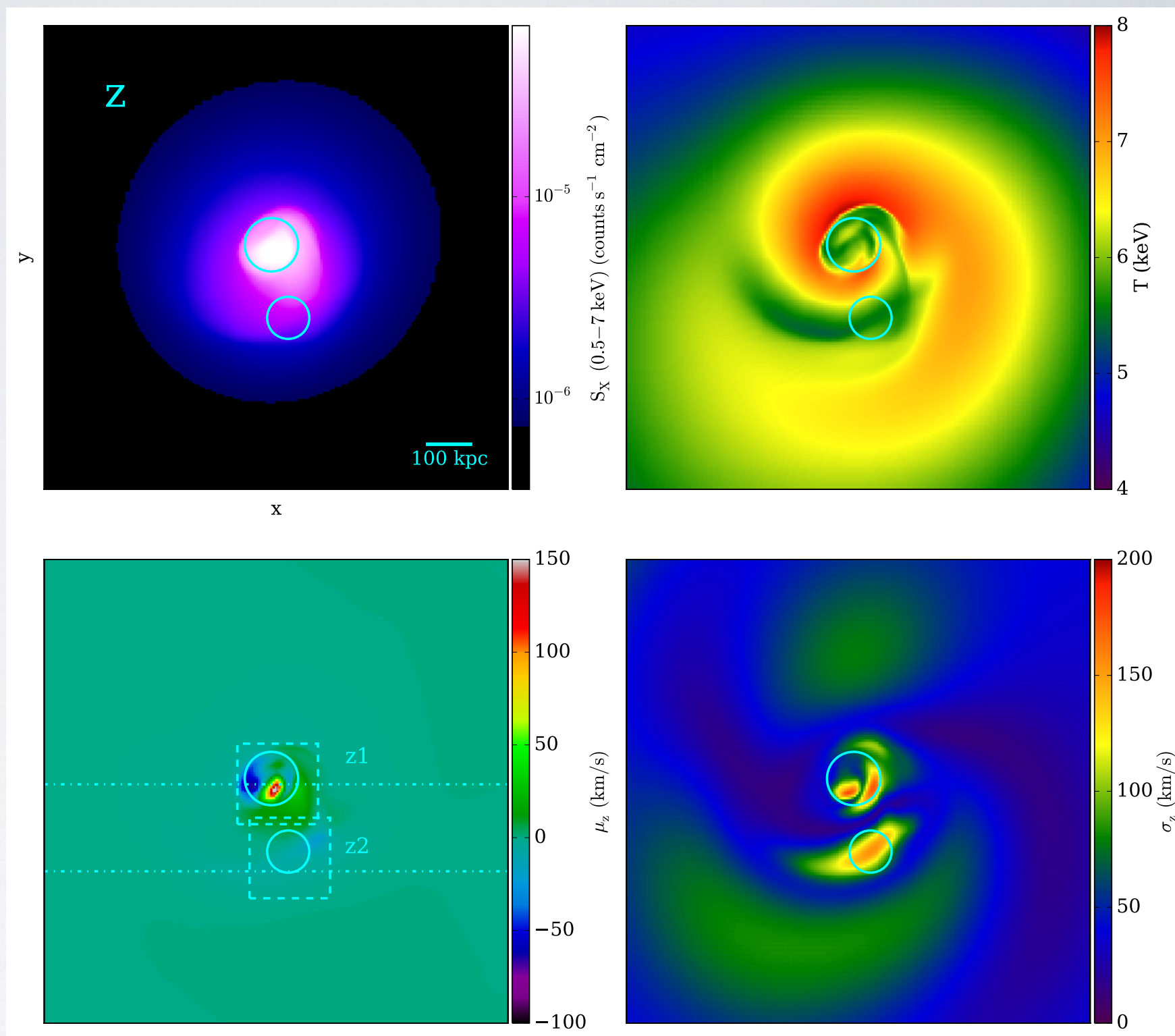
Inviscid

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LOOKING
PERP TO
PLANE



μ

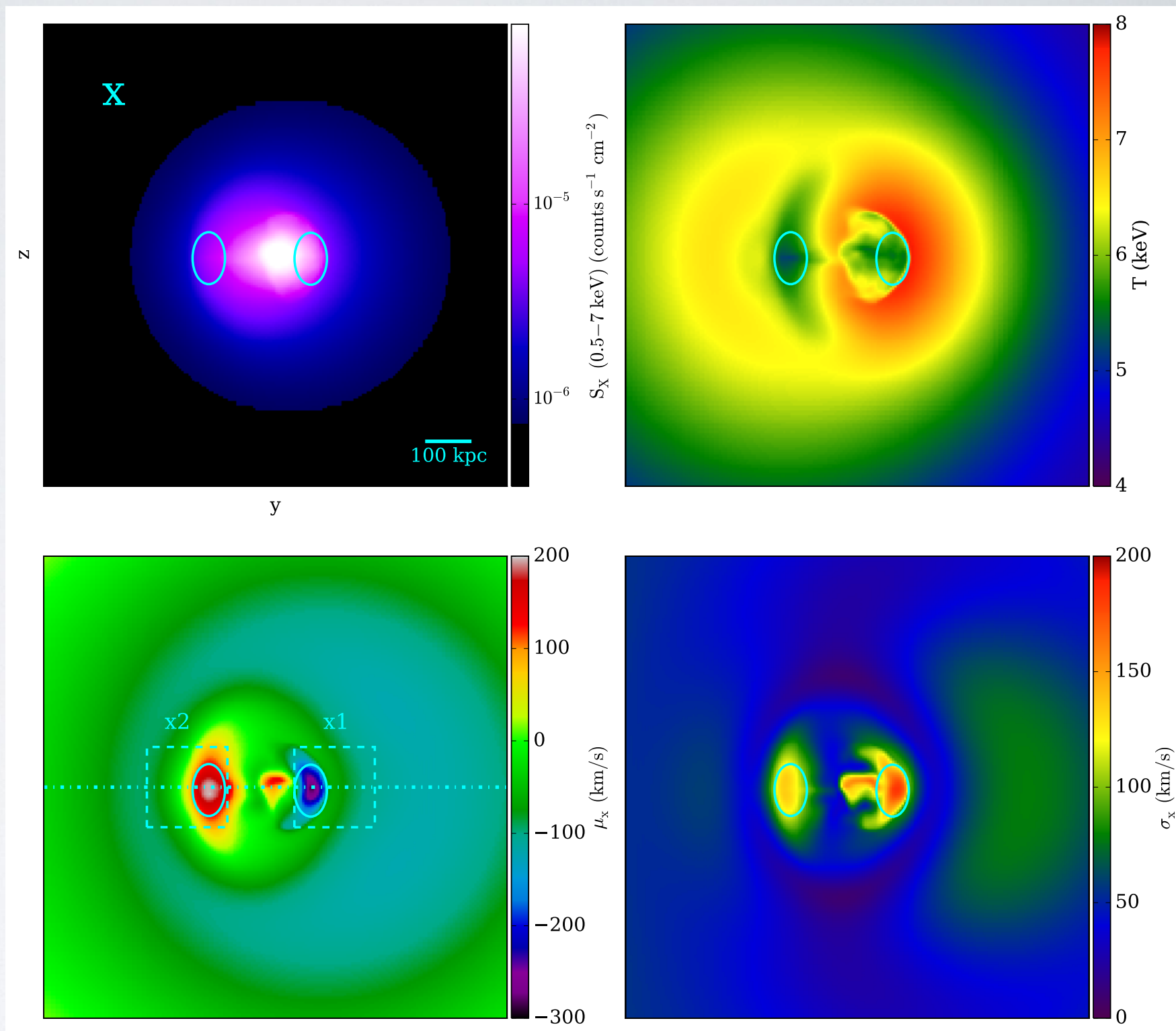
σ

Viscous

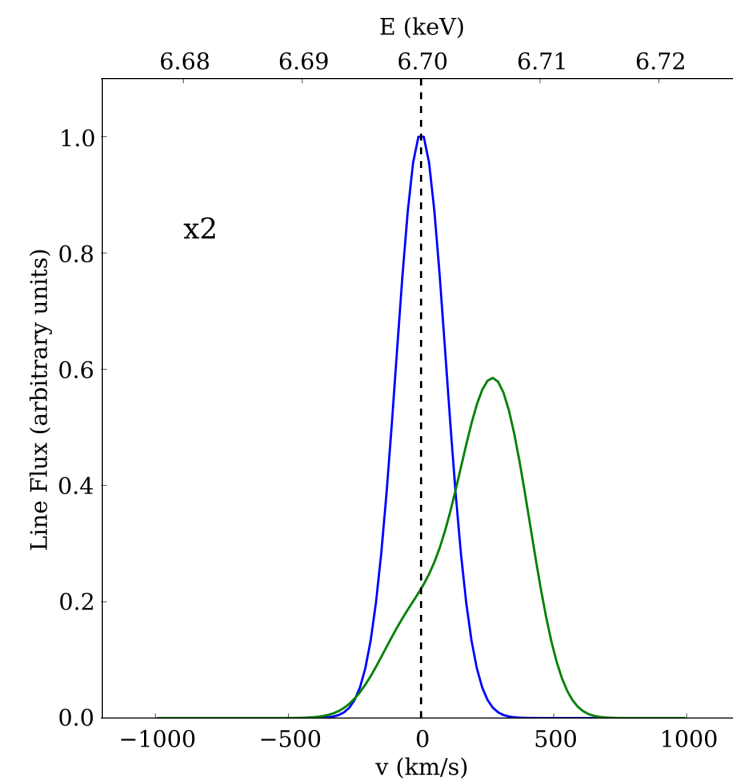
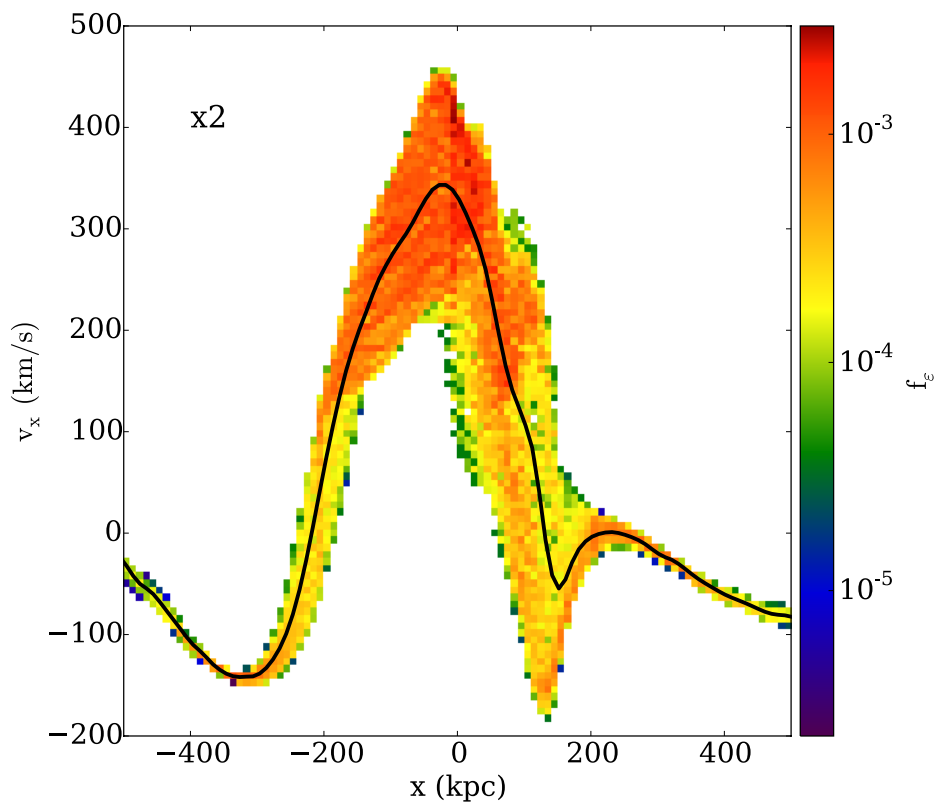
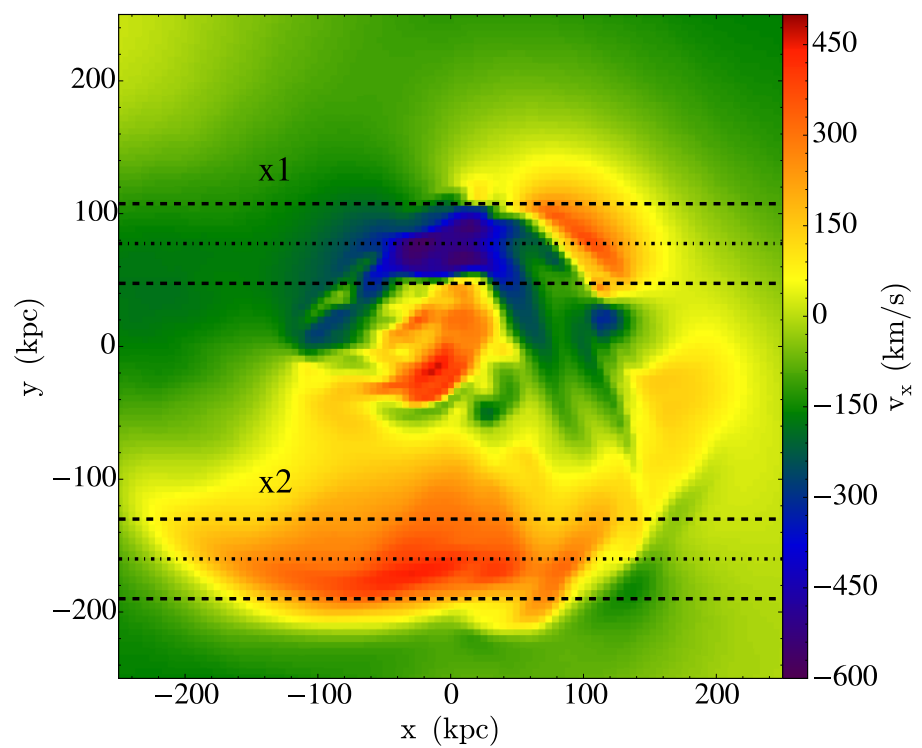
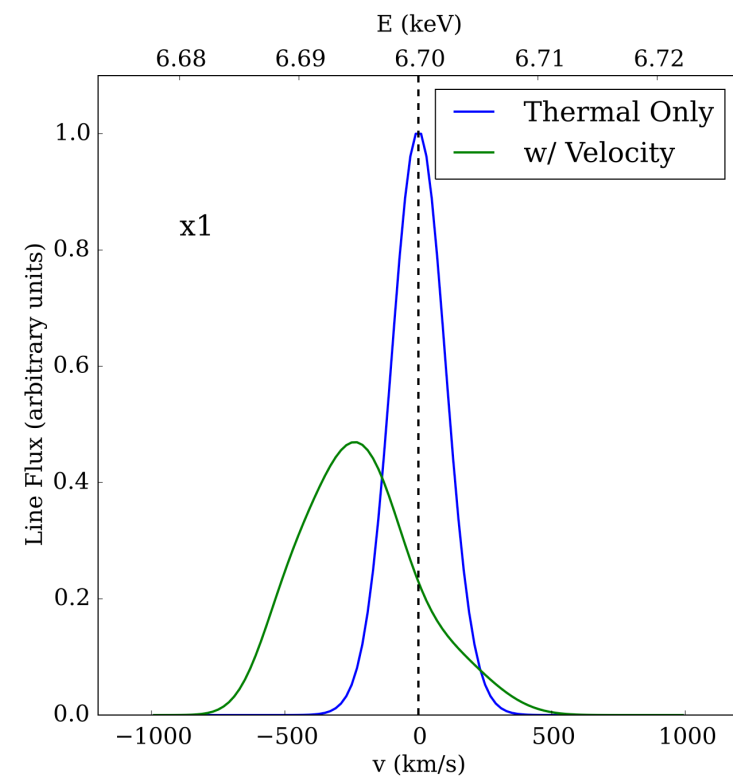
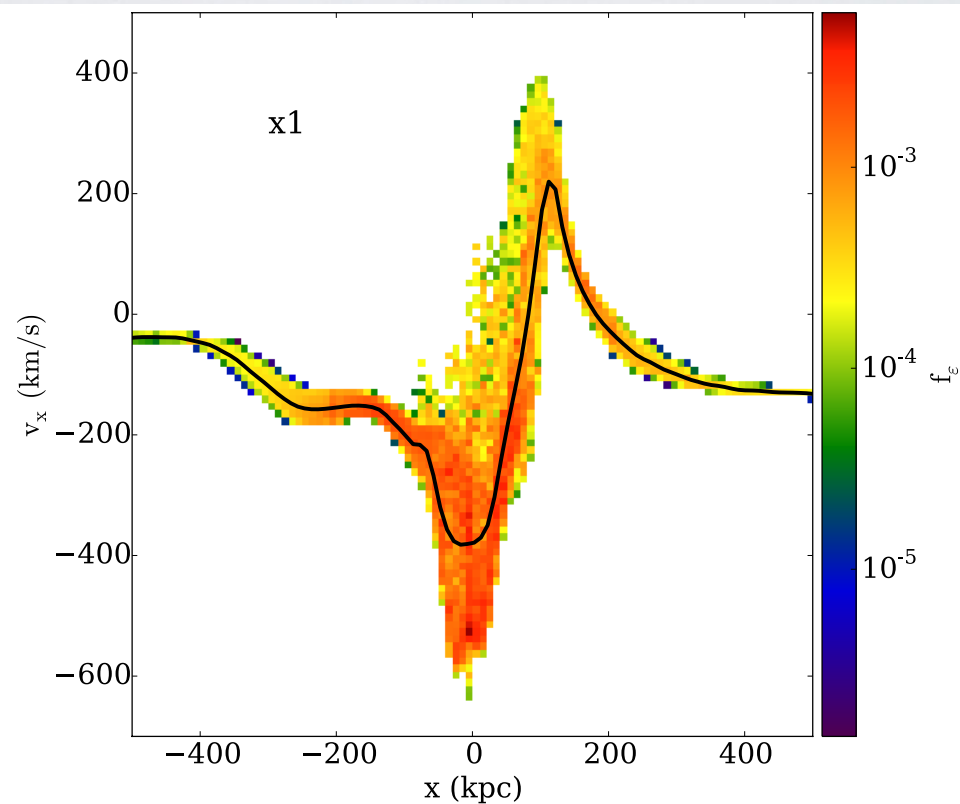
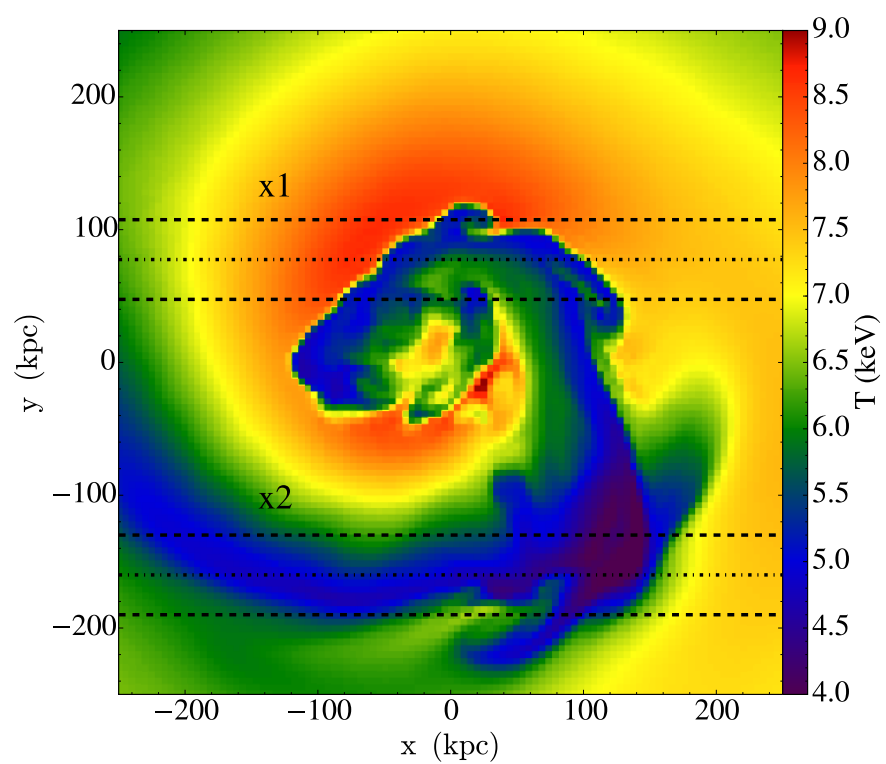
SB

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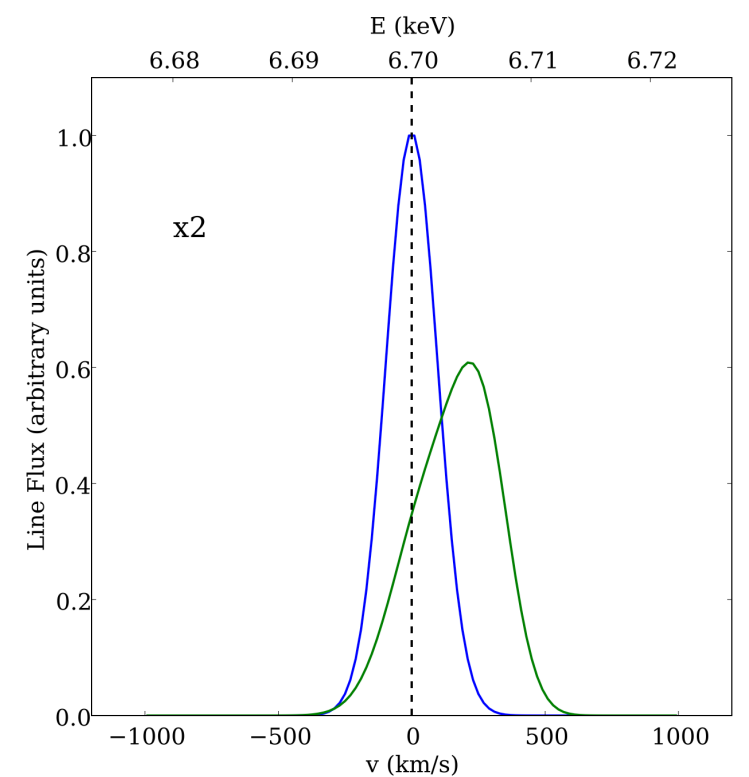
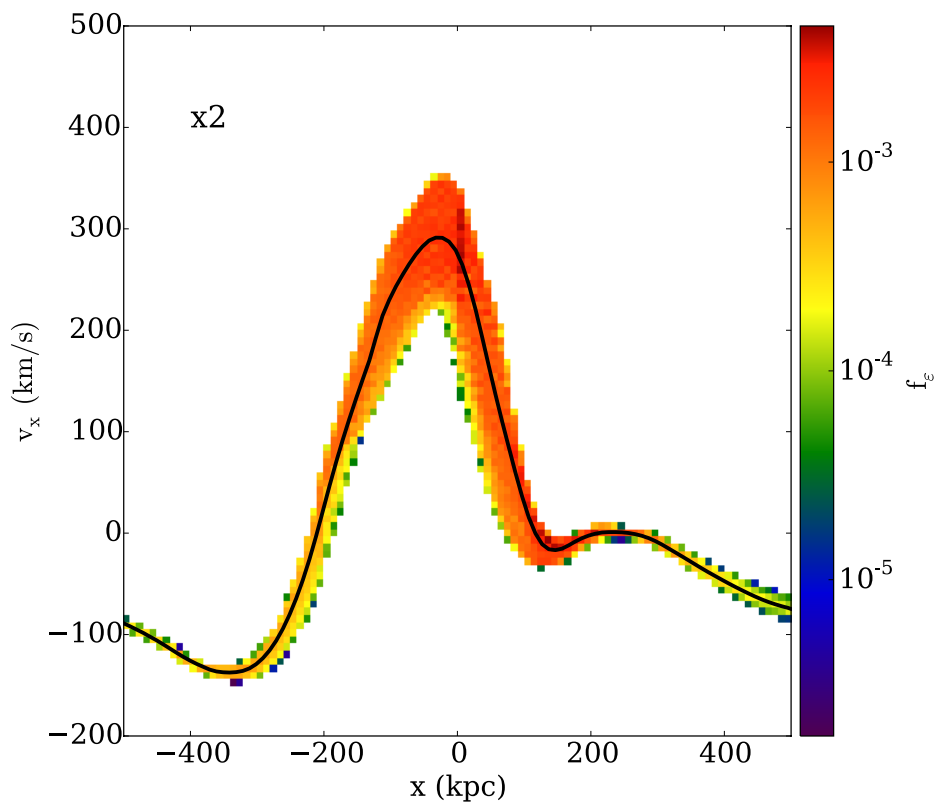
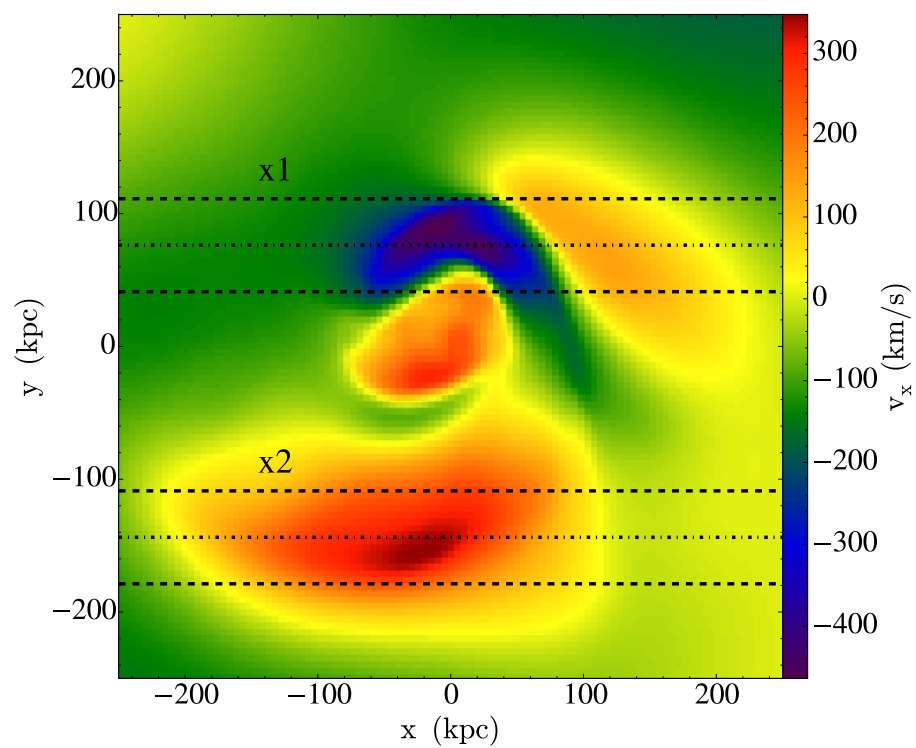
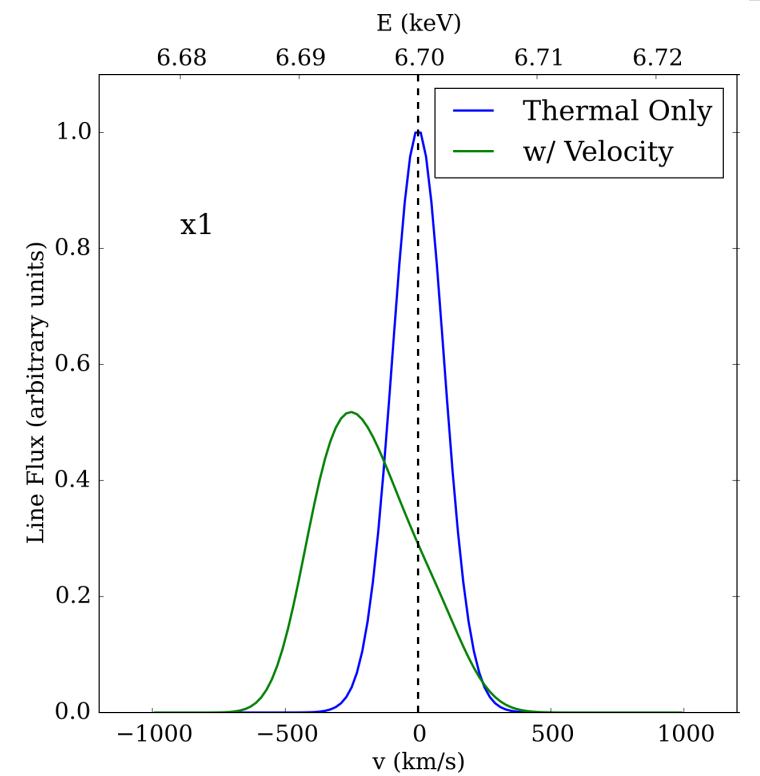
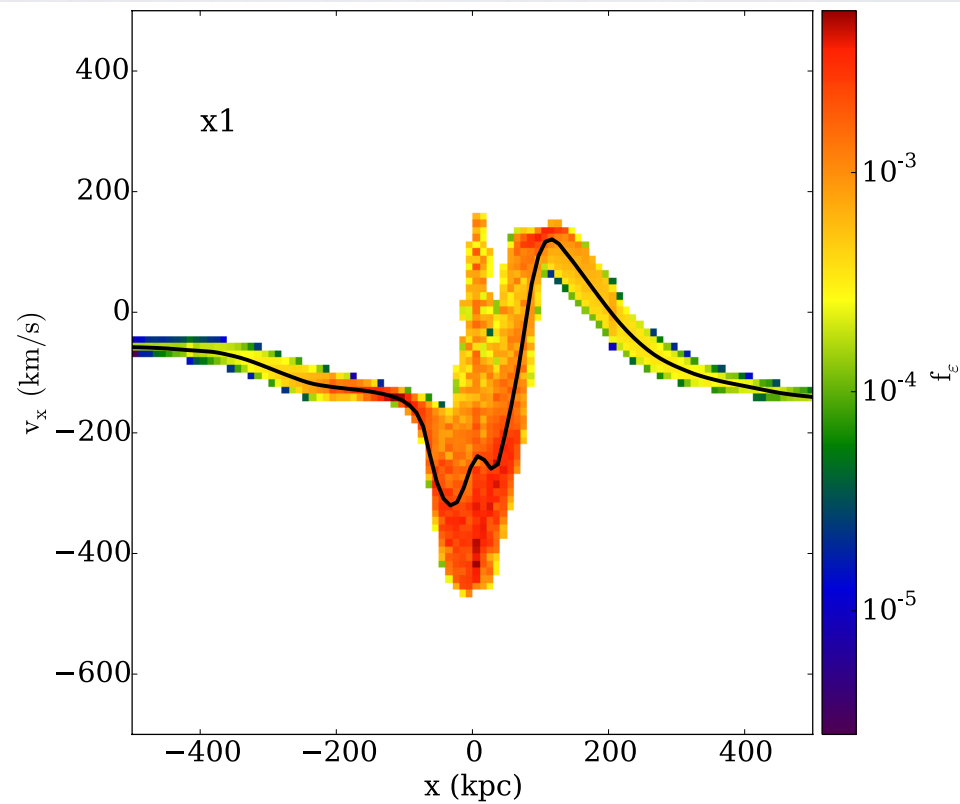
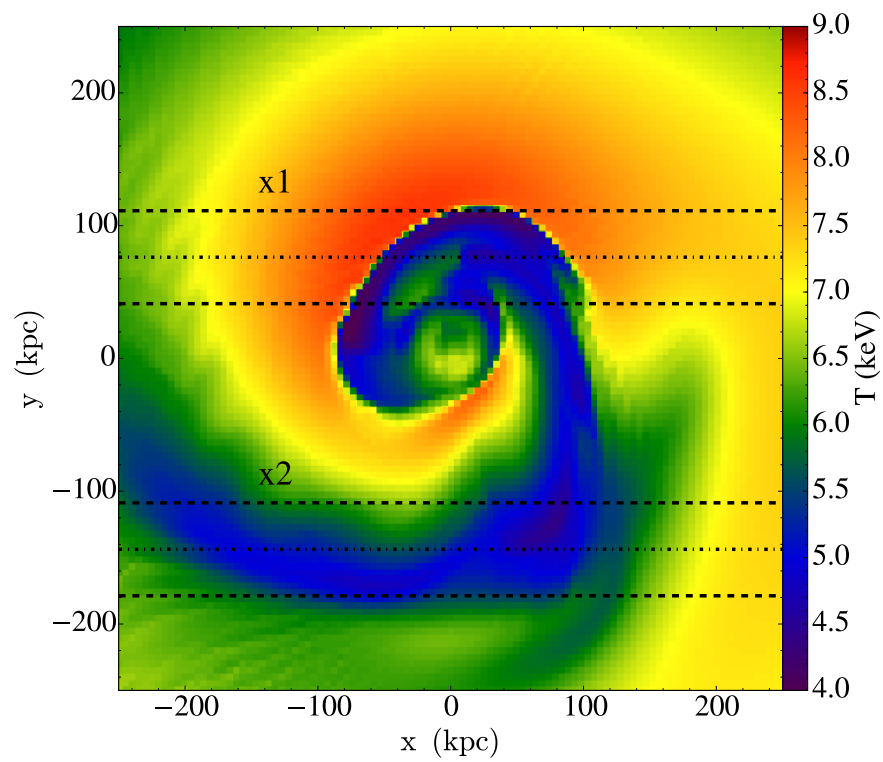
LOOKING INTO THE PLANE



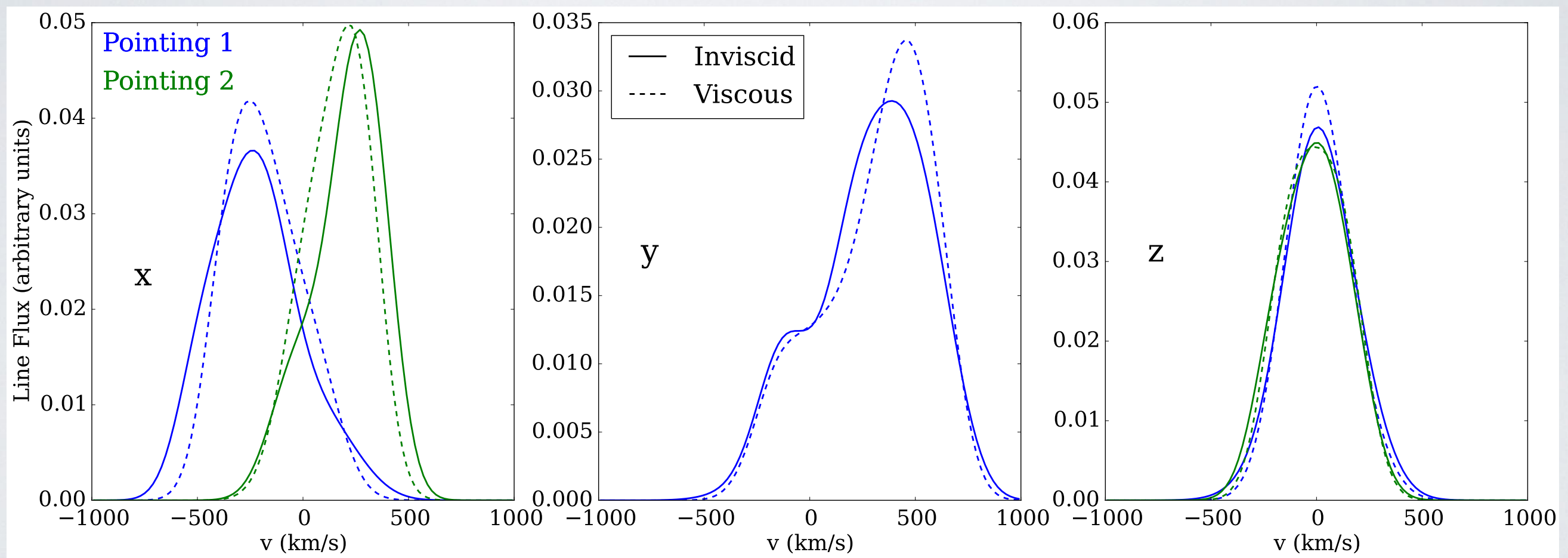
INVISCID: IN THE PLANE



VISCOUS: IN THE PLANE

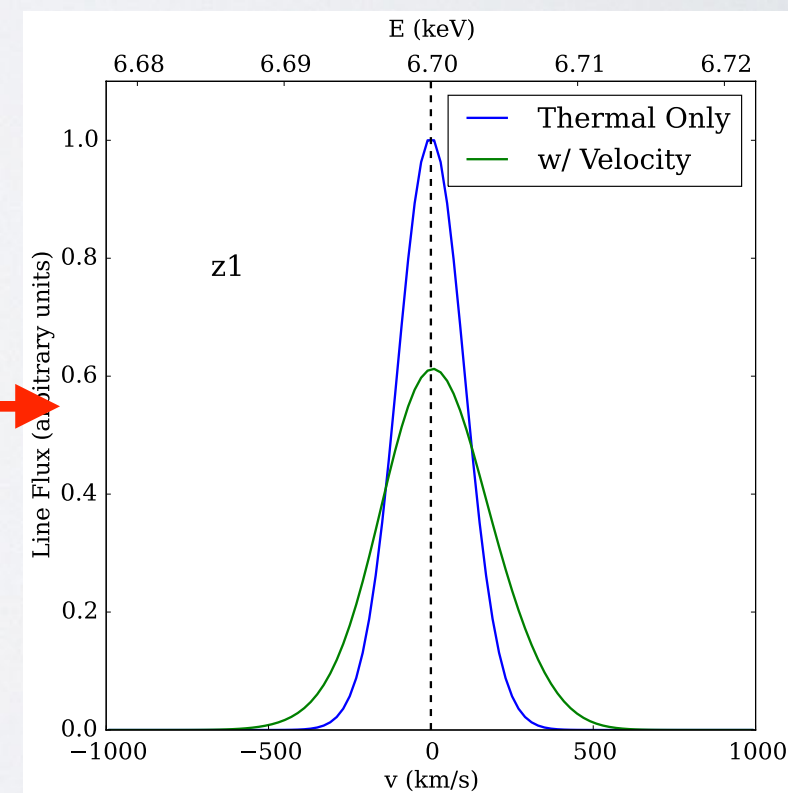
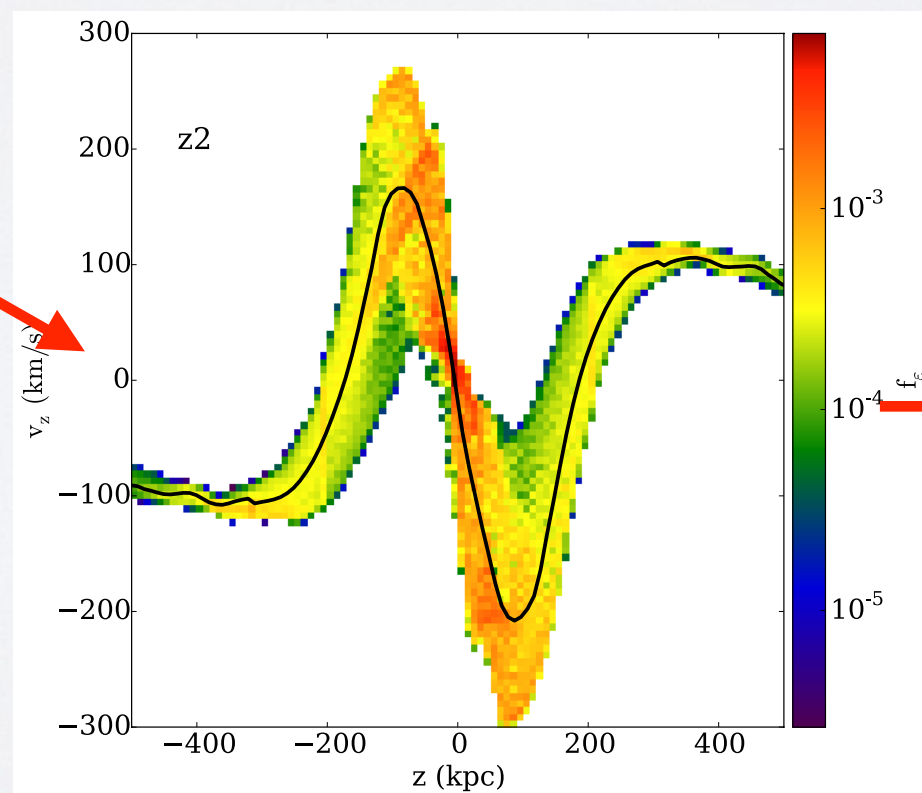
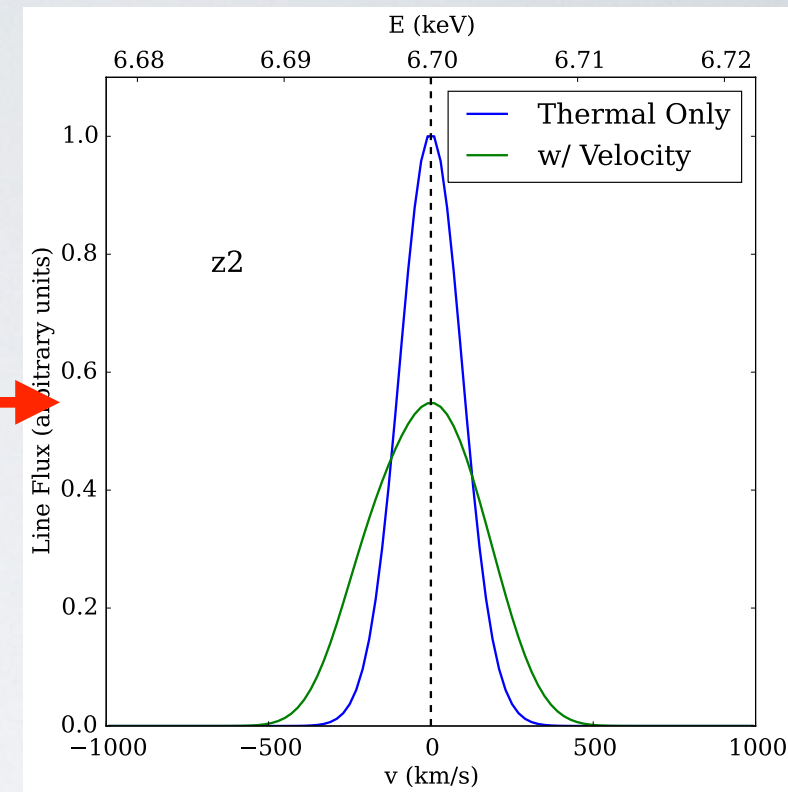
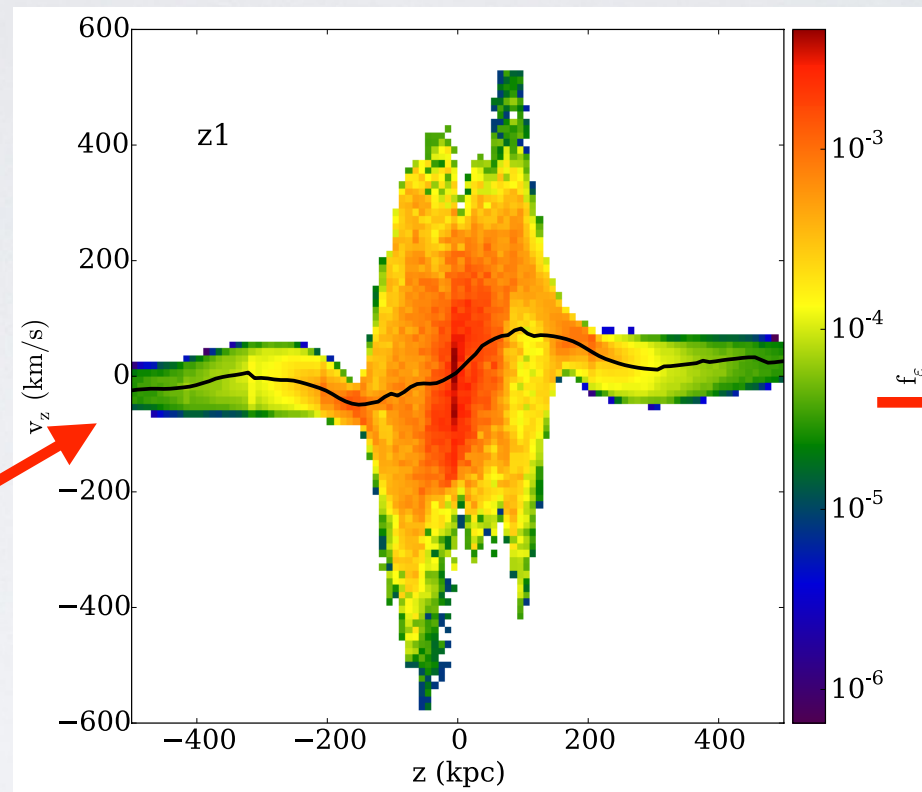
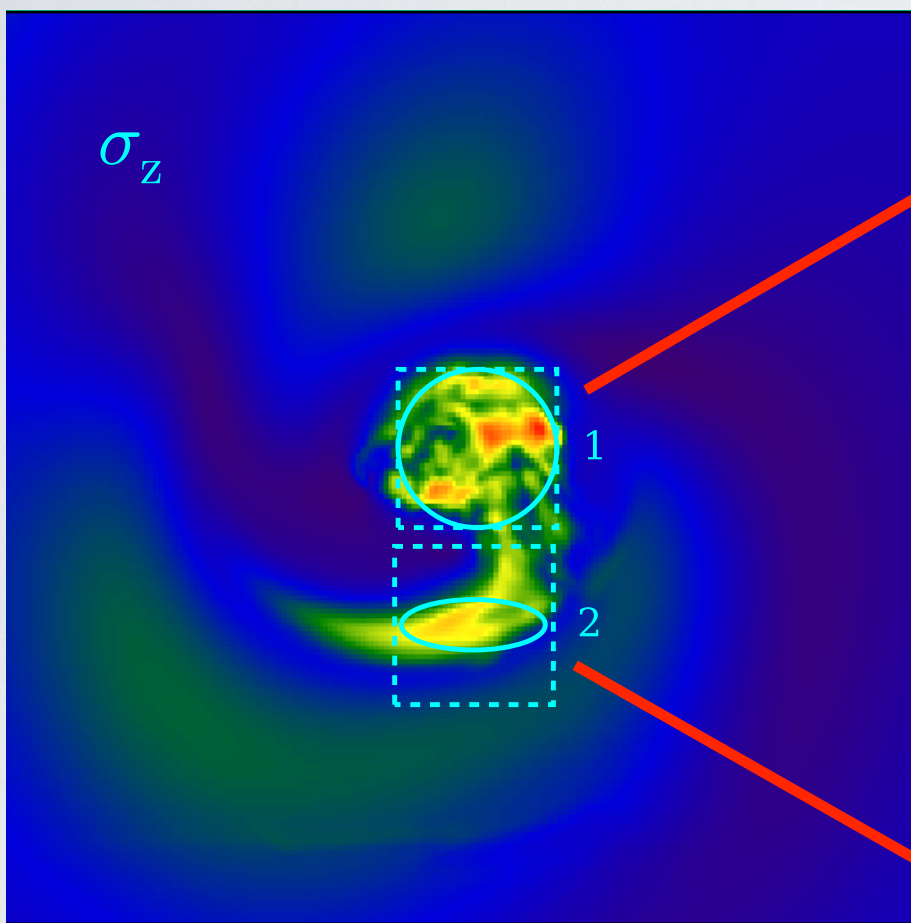


DIFFERENT VISCOSITY, SAME SHAPE

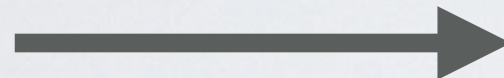
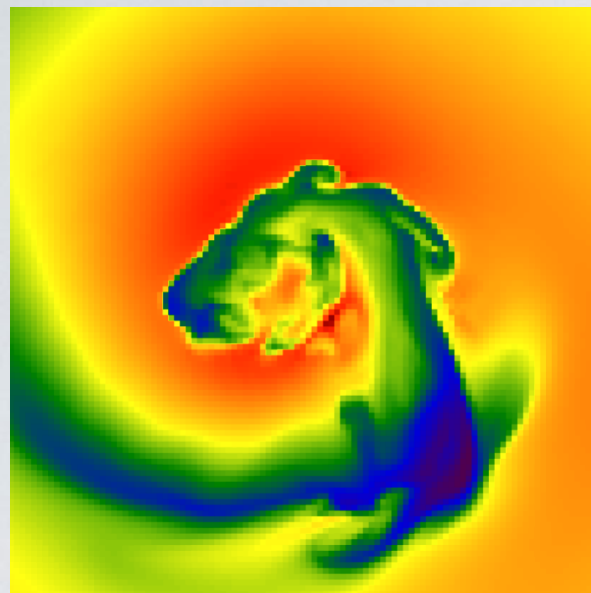


arXiv:1508.04426

DIFFERENT MOTIONS, SAME SHAPE (SOMETIMES)



SYNTHETIC OBSERVATIONS

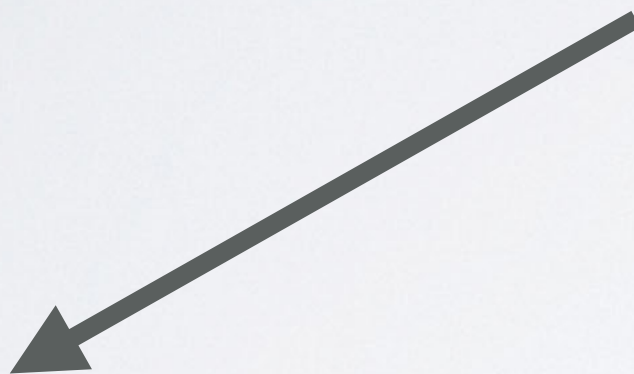


+ PHOX

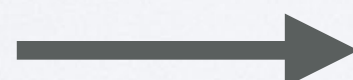
<http://yt-project.org>

<http://www.mpa-garching.mpg.de/~kdolag/Phox/>

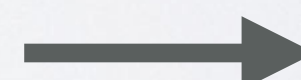
(Biffi et al 2012, 2013, MNRAS)



event lists: RA, Dec, E



SIMX



**“real”
event files**

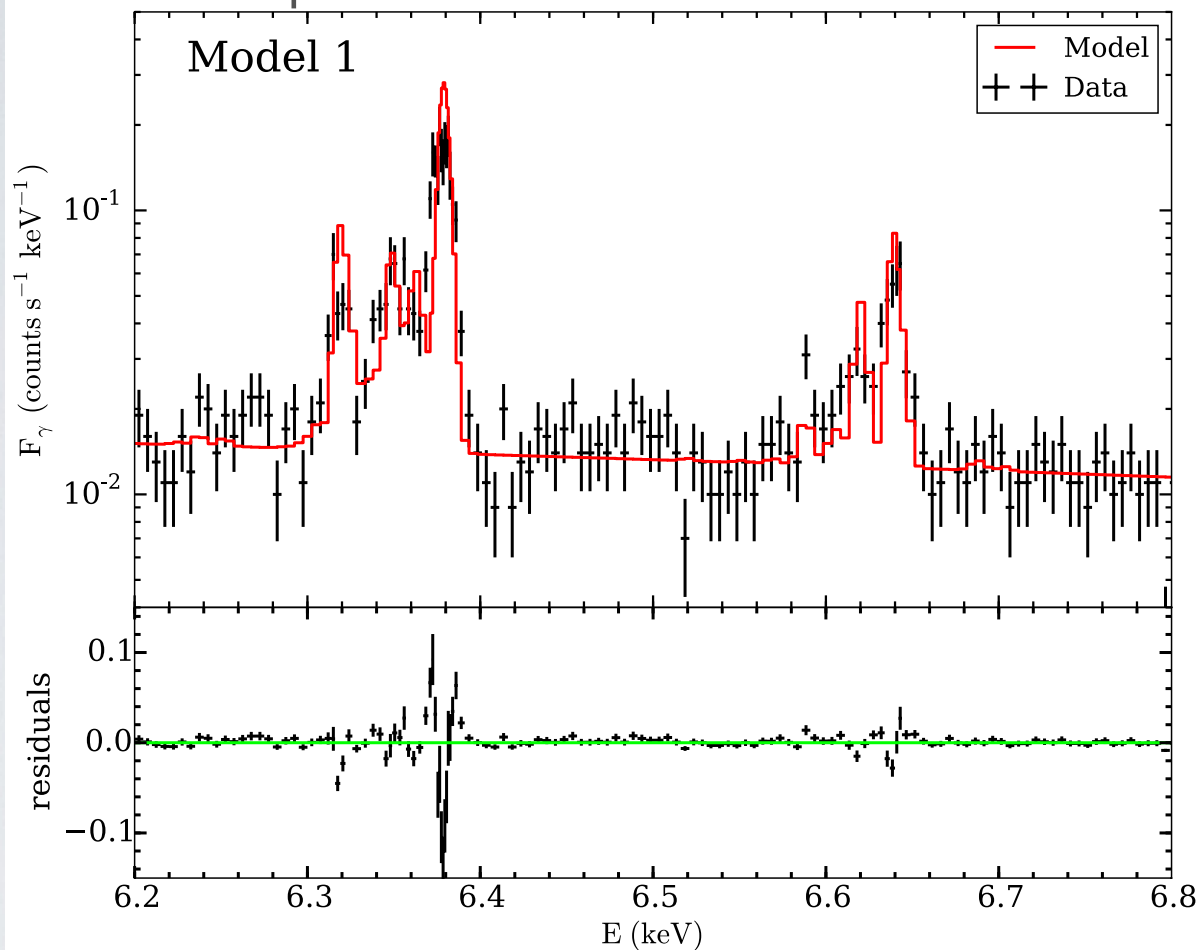
<http://hea-www.harvard.edu/simx/>

ZuHone et al 2014, arXiv:1407.1783

<http://www.youtube.com/watch?v=fUMq6rmNshc>

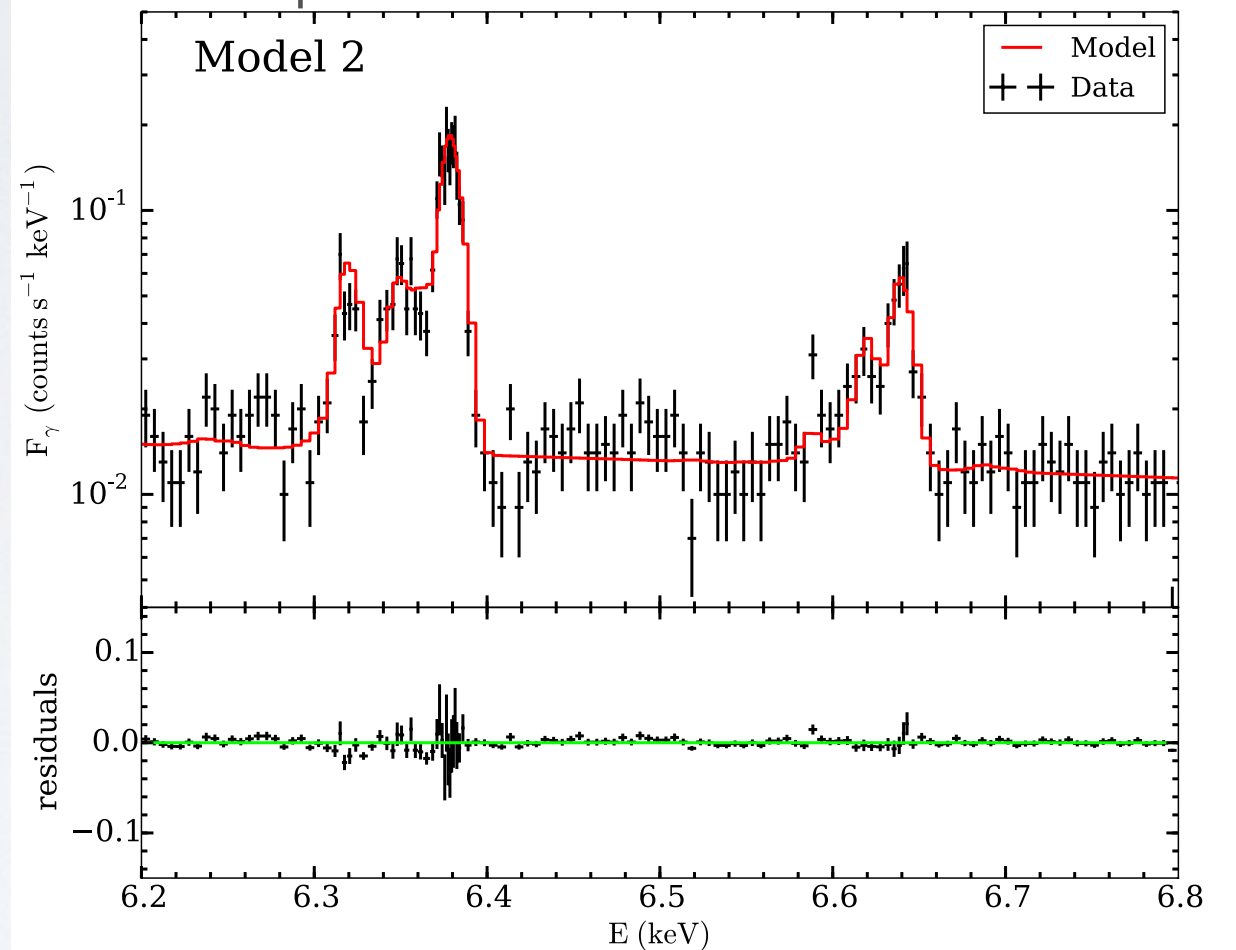
FITTING SPECTRA

exposure time = 200 ks

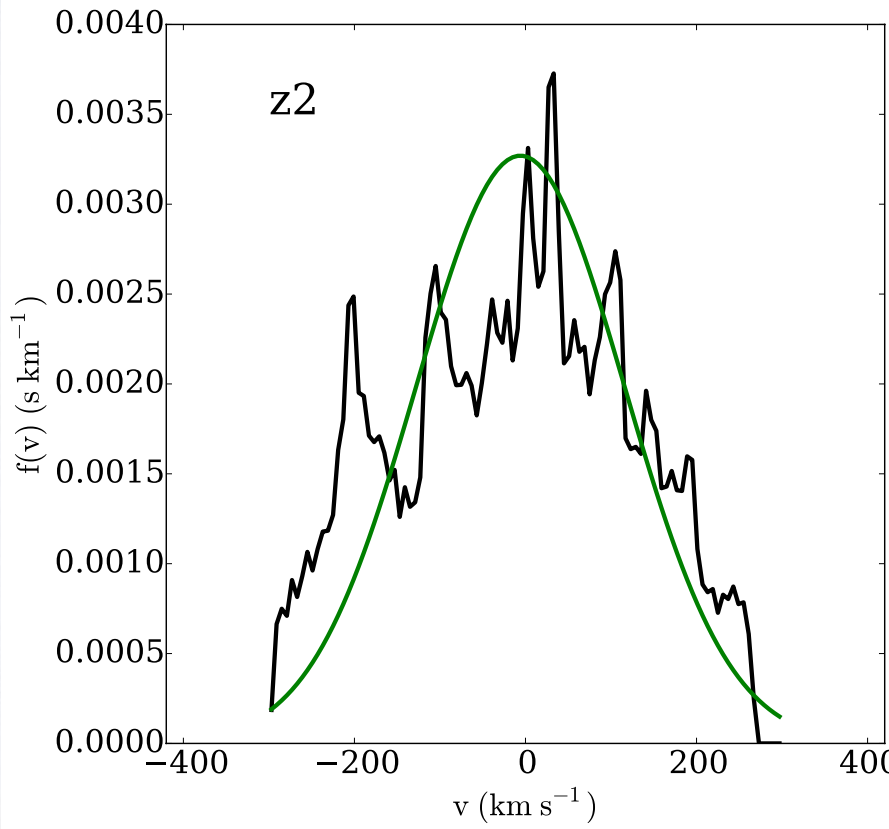
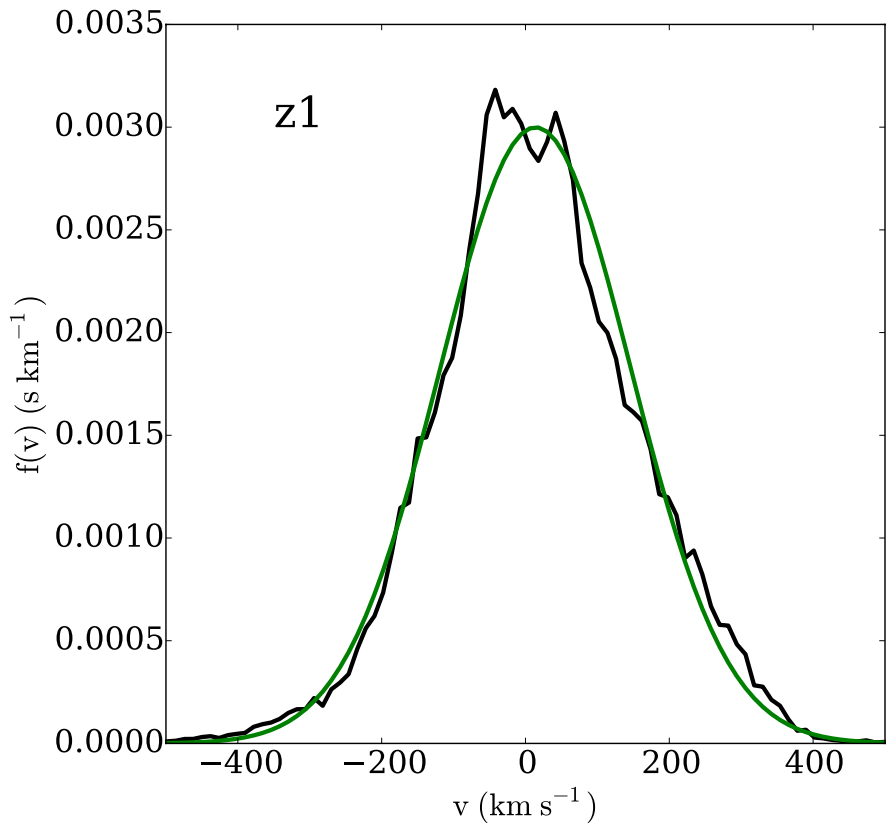
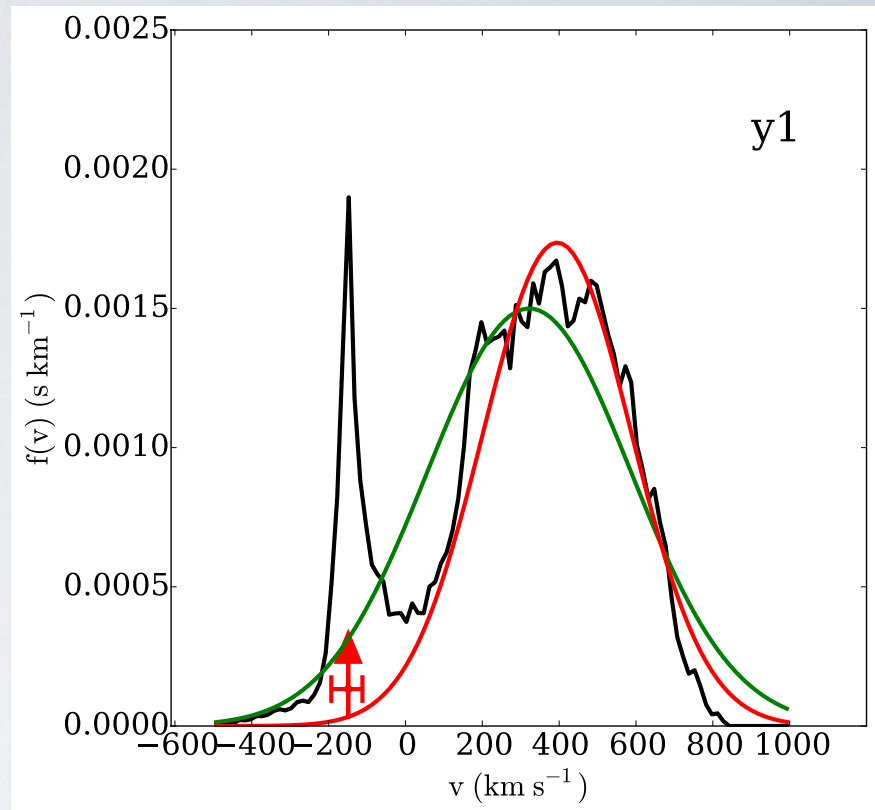
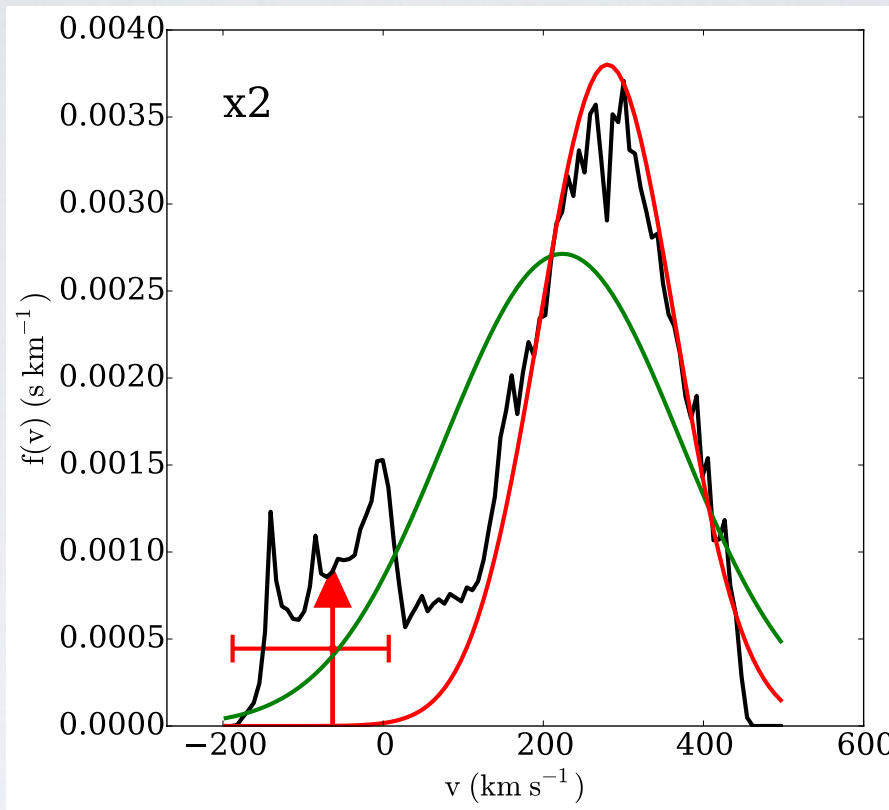
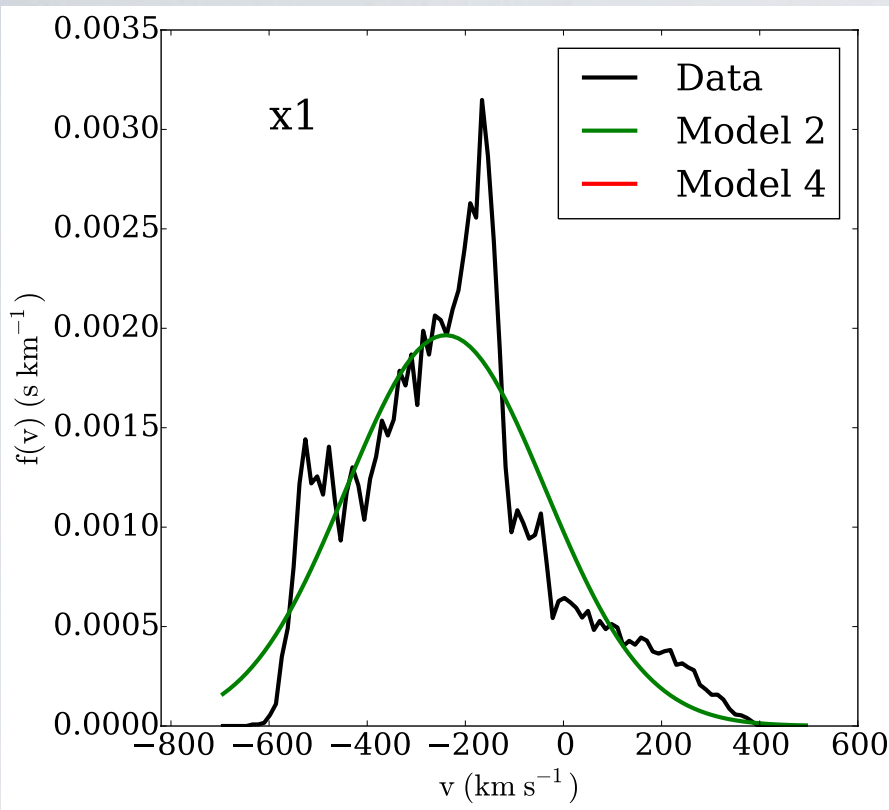


Model 1: Single-T APEC
model, w/ thermal
broadening only

exposure time = 200 ks



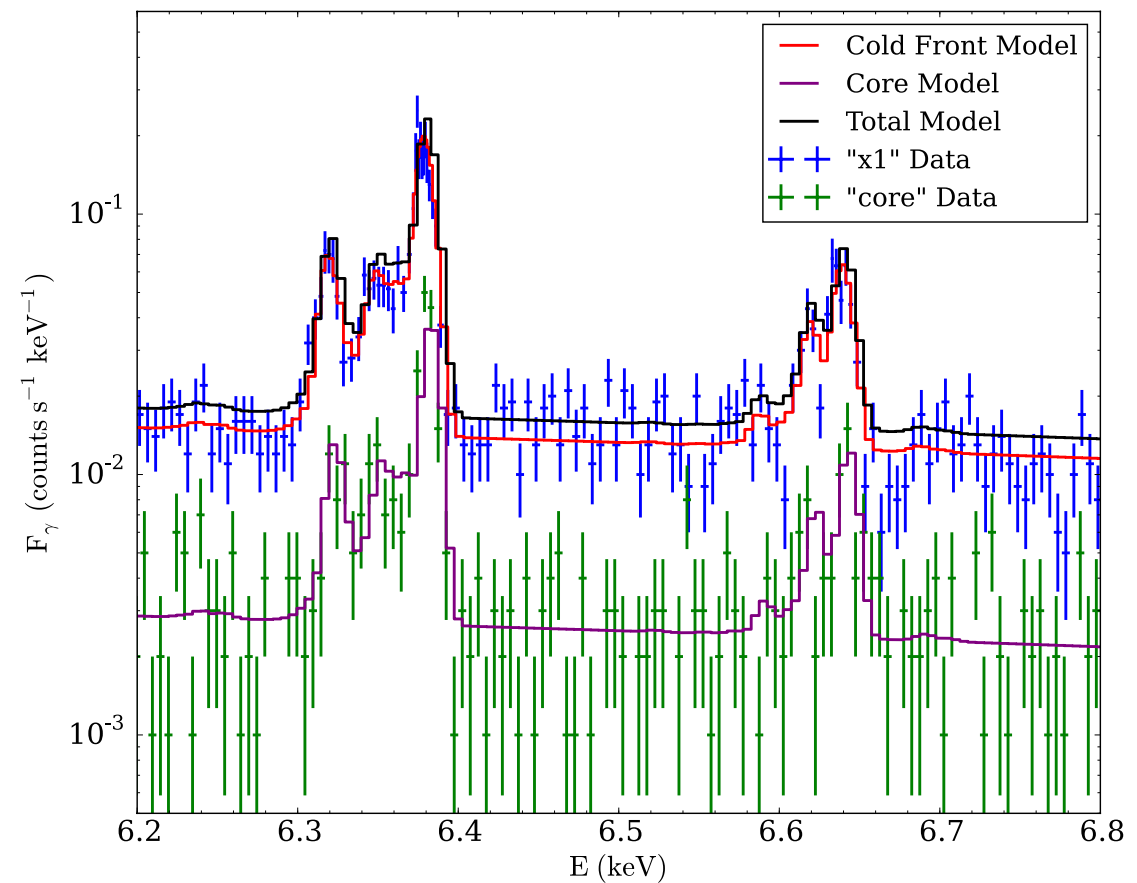
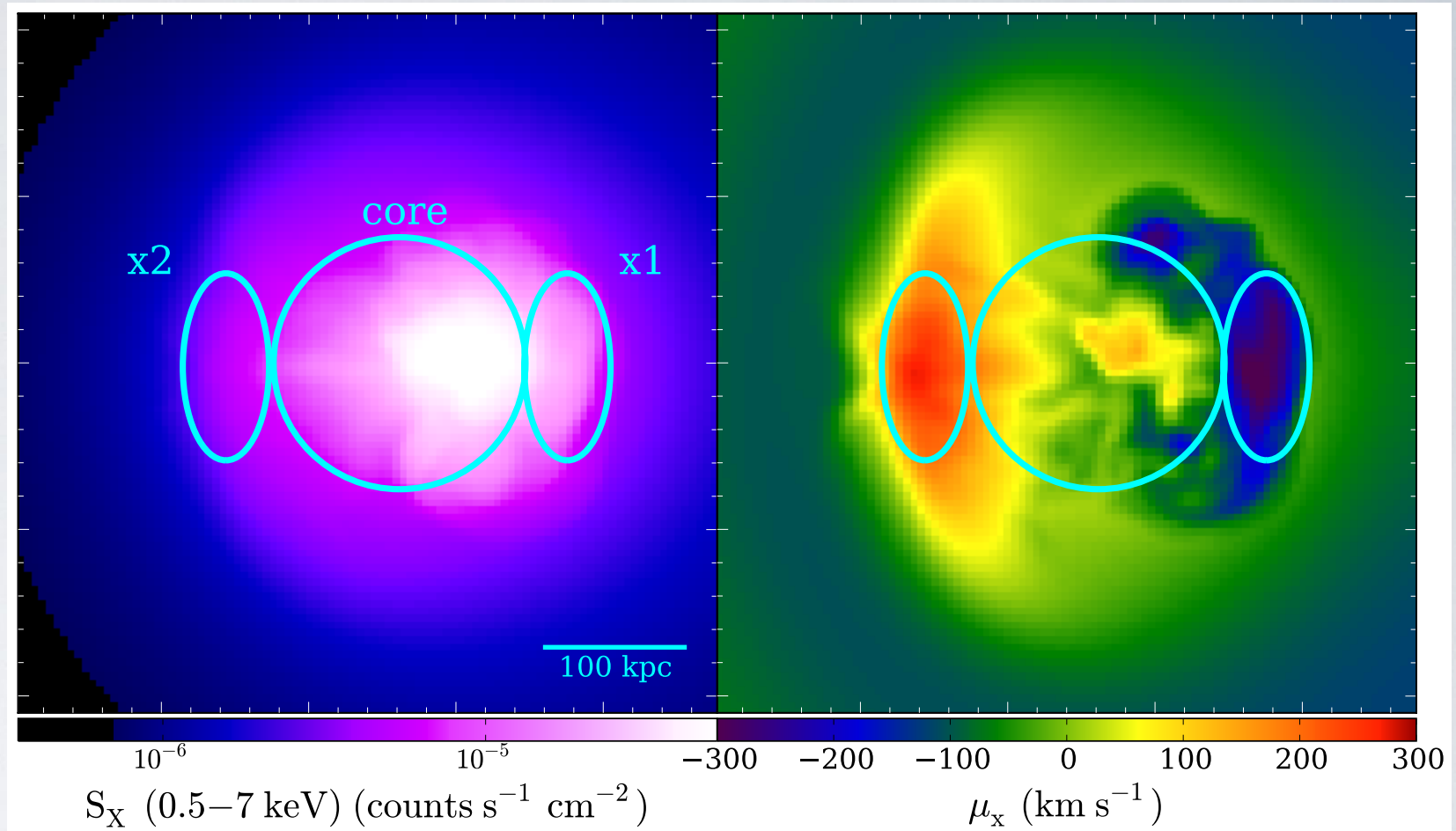
Model 2: Single-T APEC
model, w/ thermal and
Gaussian velocity
broadening



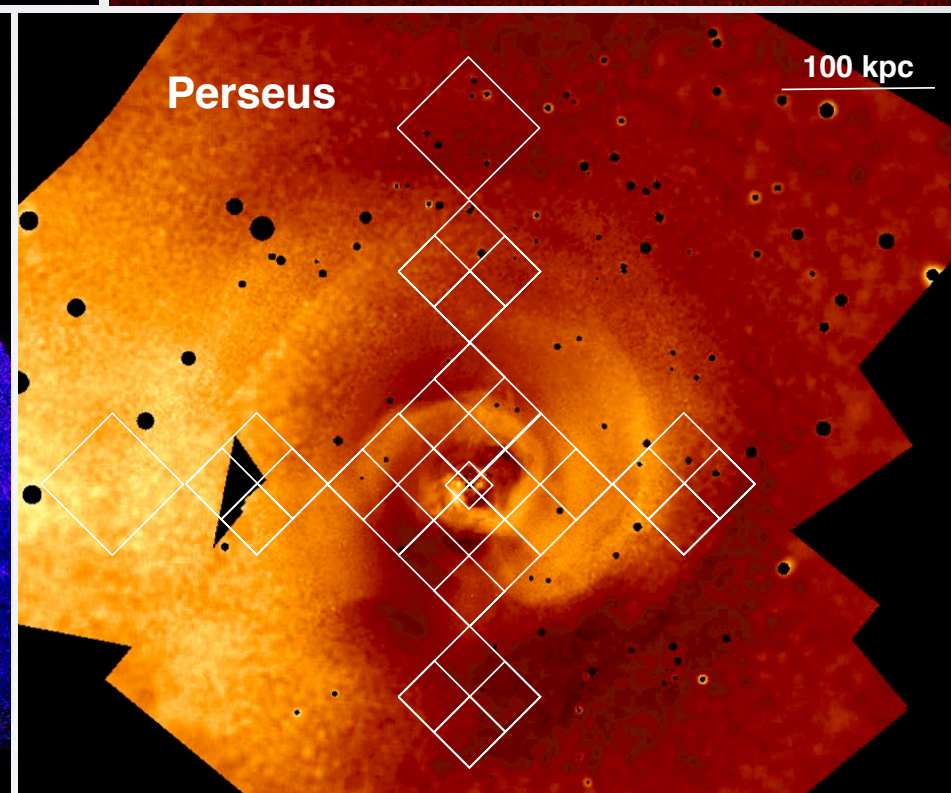
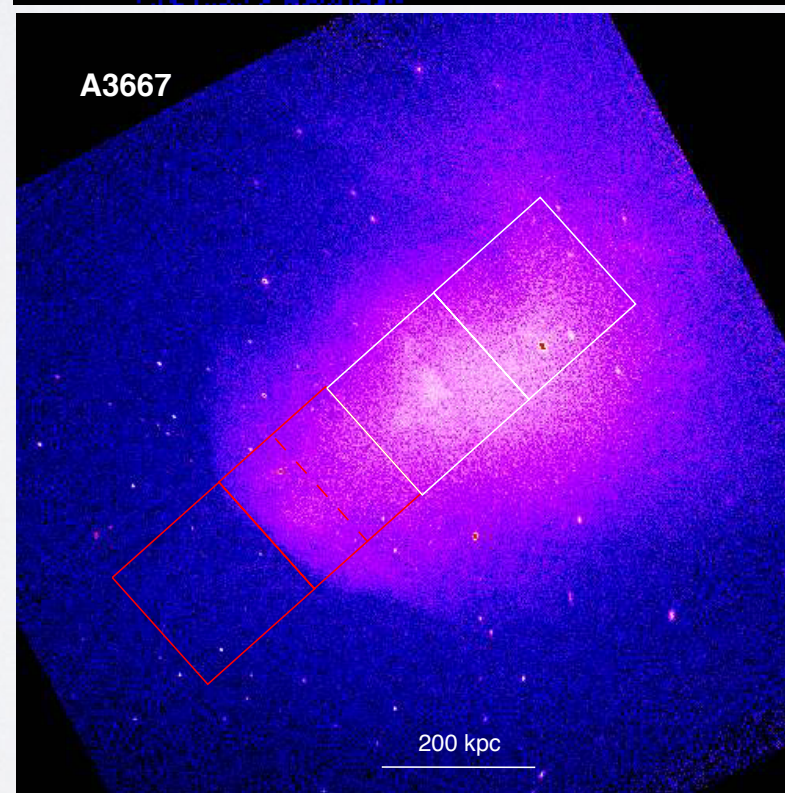
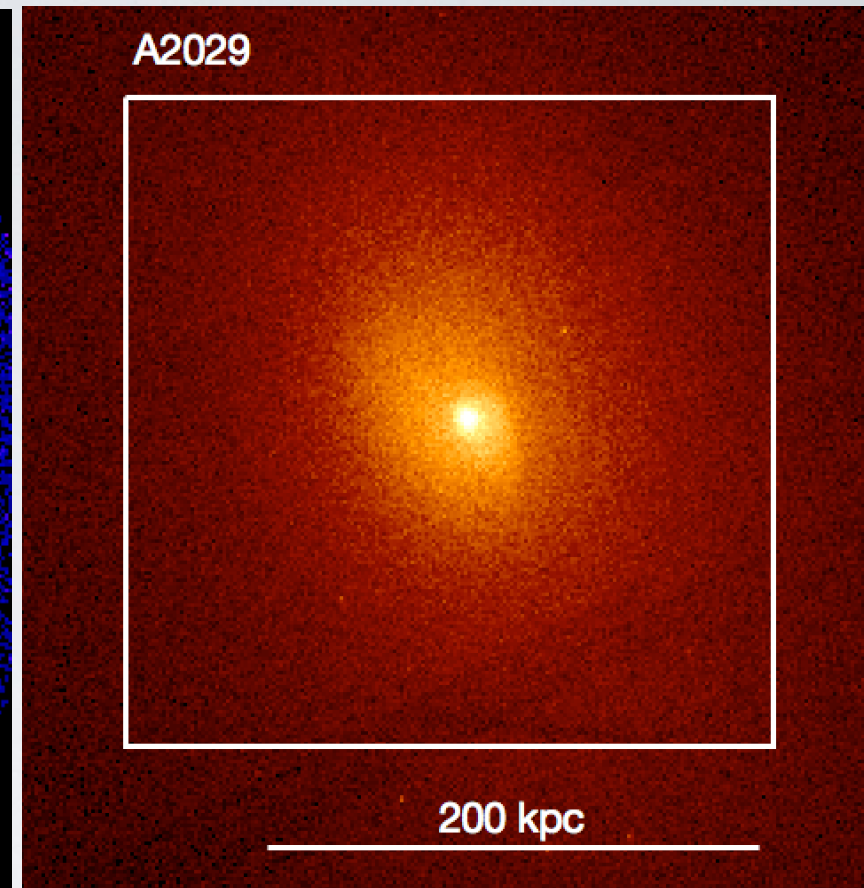
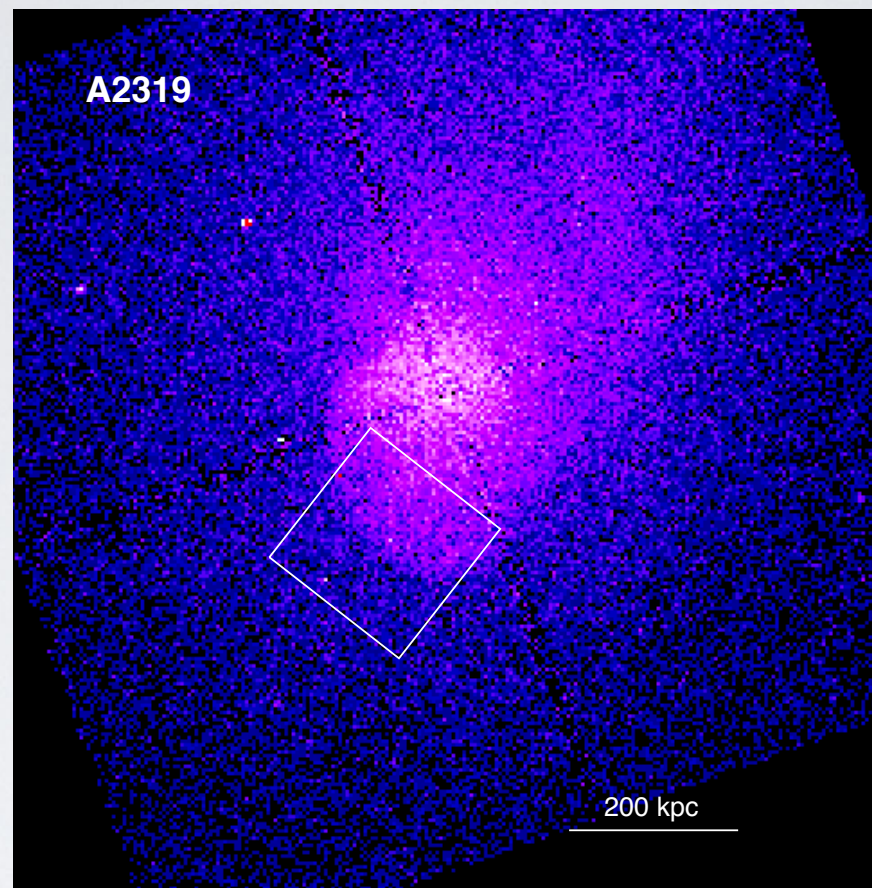
Predicted vs. Real
Velocity
Distributions

ACCOUNTING FOR PSF SCATTERING

[arXiv:1508.04426](https://arxiv.org/abs/1508.04426)



Where
should
we look?



Astro-H →

Athena →

Inviscid

Shift

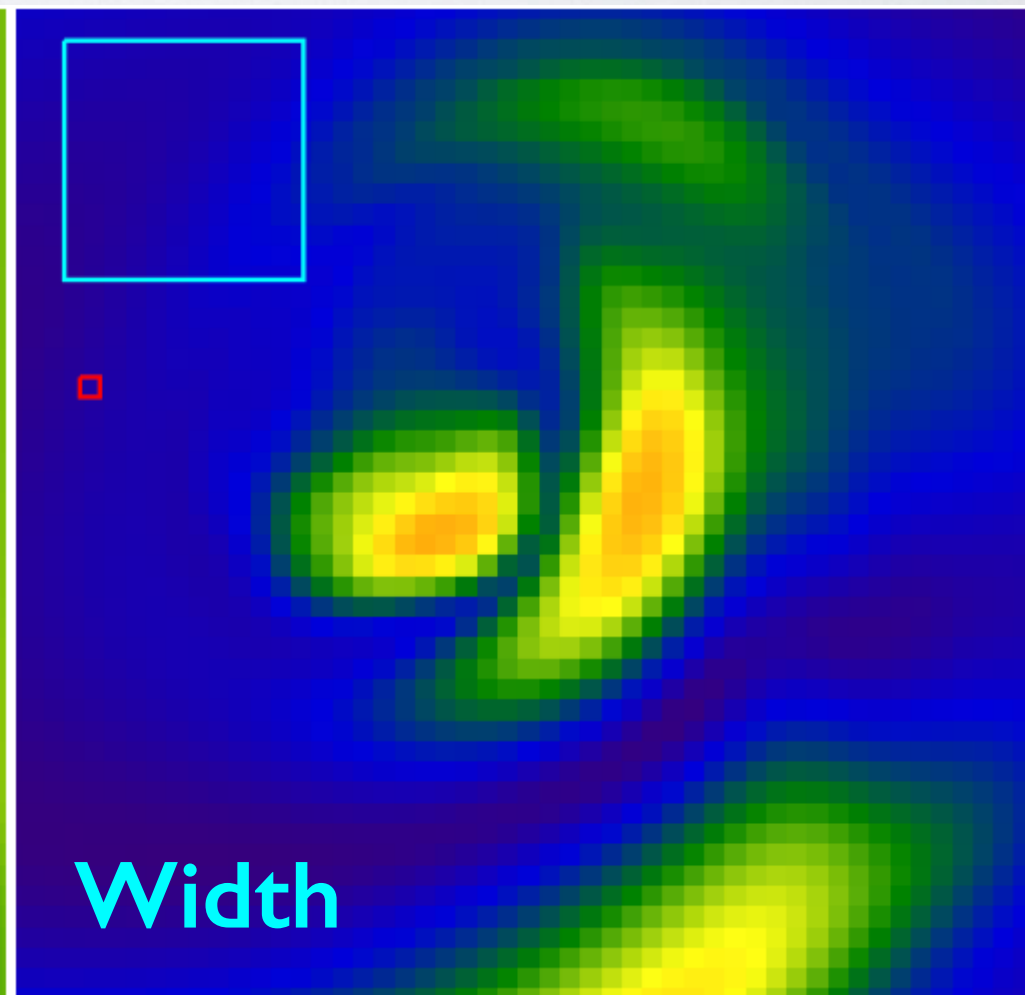
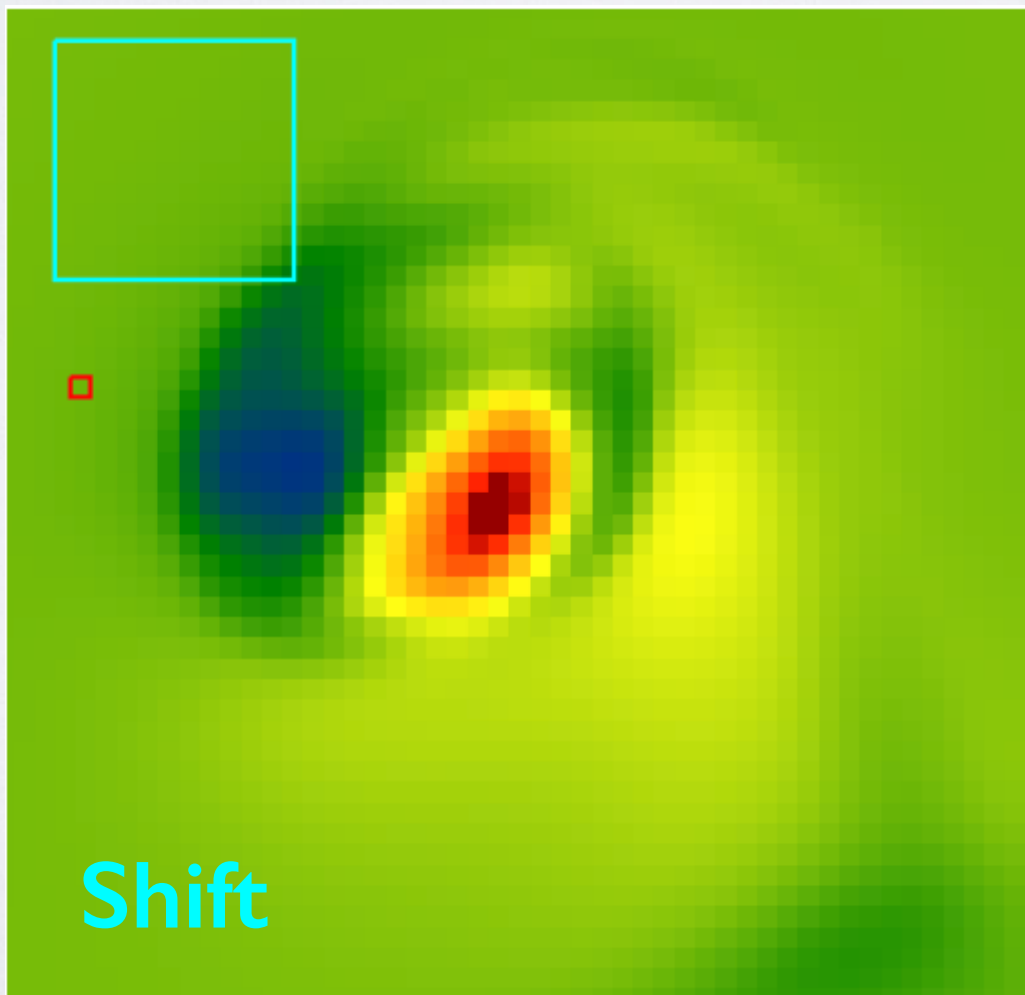
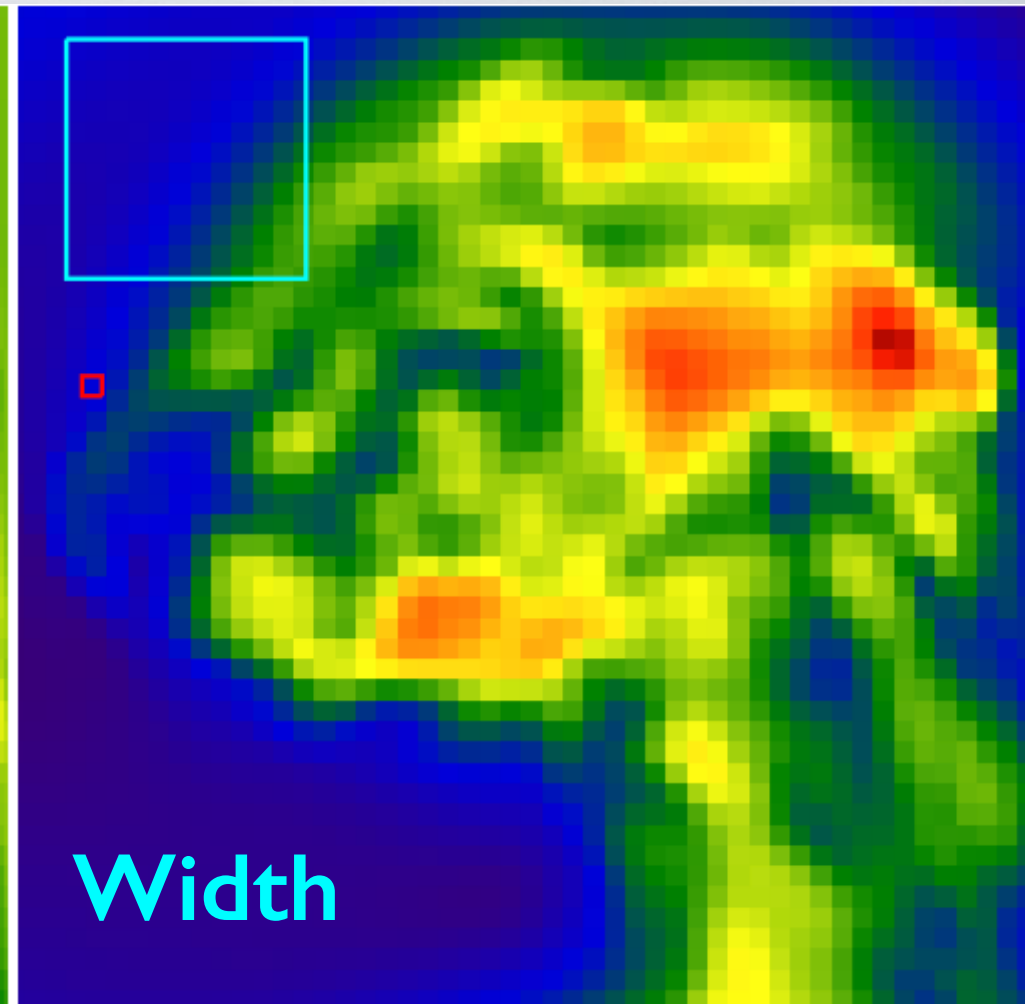
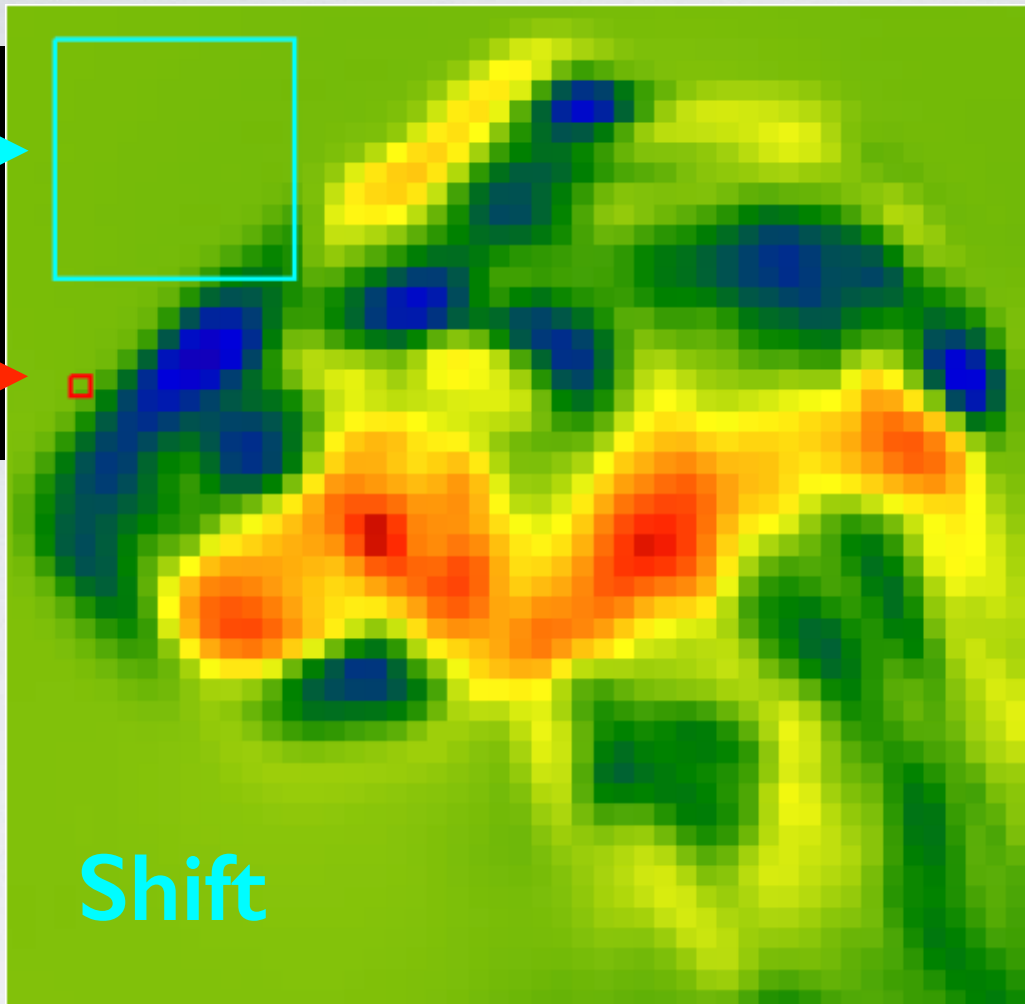
Width

Cluster at
 $z = 0.05$

Viscous

Shift

Width



SUMMARY

- *Astro-H* will be able to measure line shifts and widths from sloshing motions in nearby galaxy clusters
- It will be difficult to distinguished line broadening from sloshing motions vs. turbulent motions
- Constraints on viscosity are probably limited
- Must worry about systematics—Gaussian models for non-Gaussian velocity distributions, PSF scattering from core
- What do we need? Better spatial resolution! (Athena, X-ray Surveyor, etc.)